

**CONNECTING THE SOLAR DYNAMO
BELOW THE SURFACE
WITH EJECTION OF TWISTED MAGNETIC
FIELDS ABOVE THE SURFACE**

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**MAX PLANCK INSTITUTE
FOR SOLAR SYSTEM RESEARCH**



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PETRI J. KÄPYLÄ, AALTO UNIVERSITY
MAARIT J. KÄPYLÄ, AALTO UNIVERSITY**

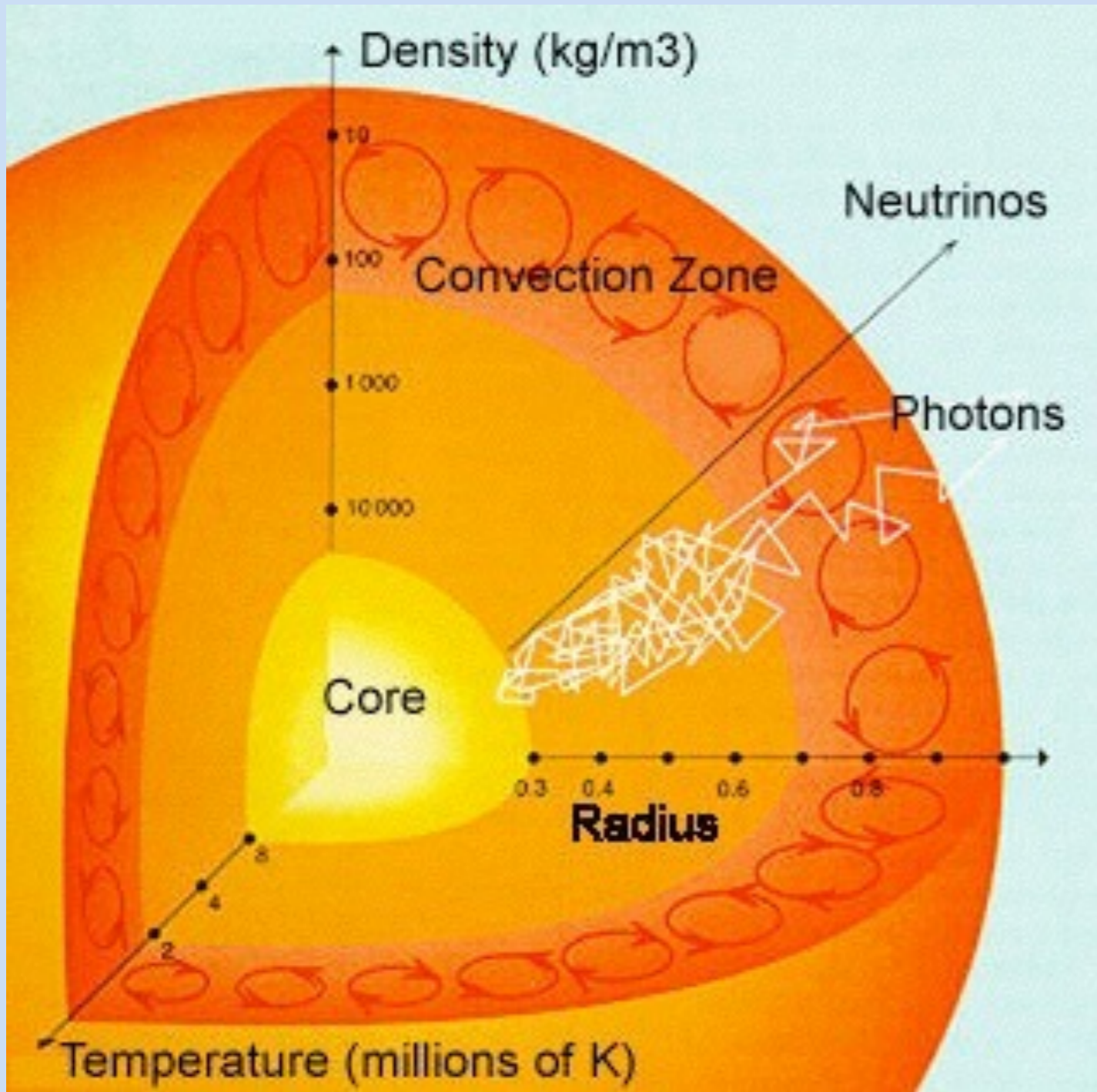


Helicity



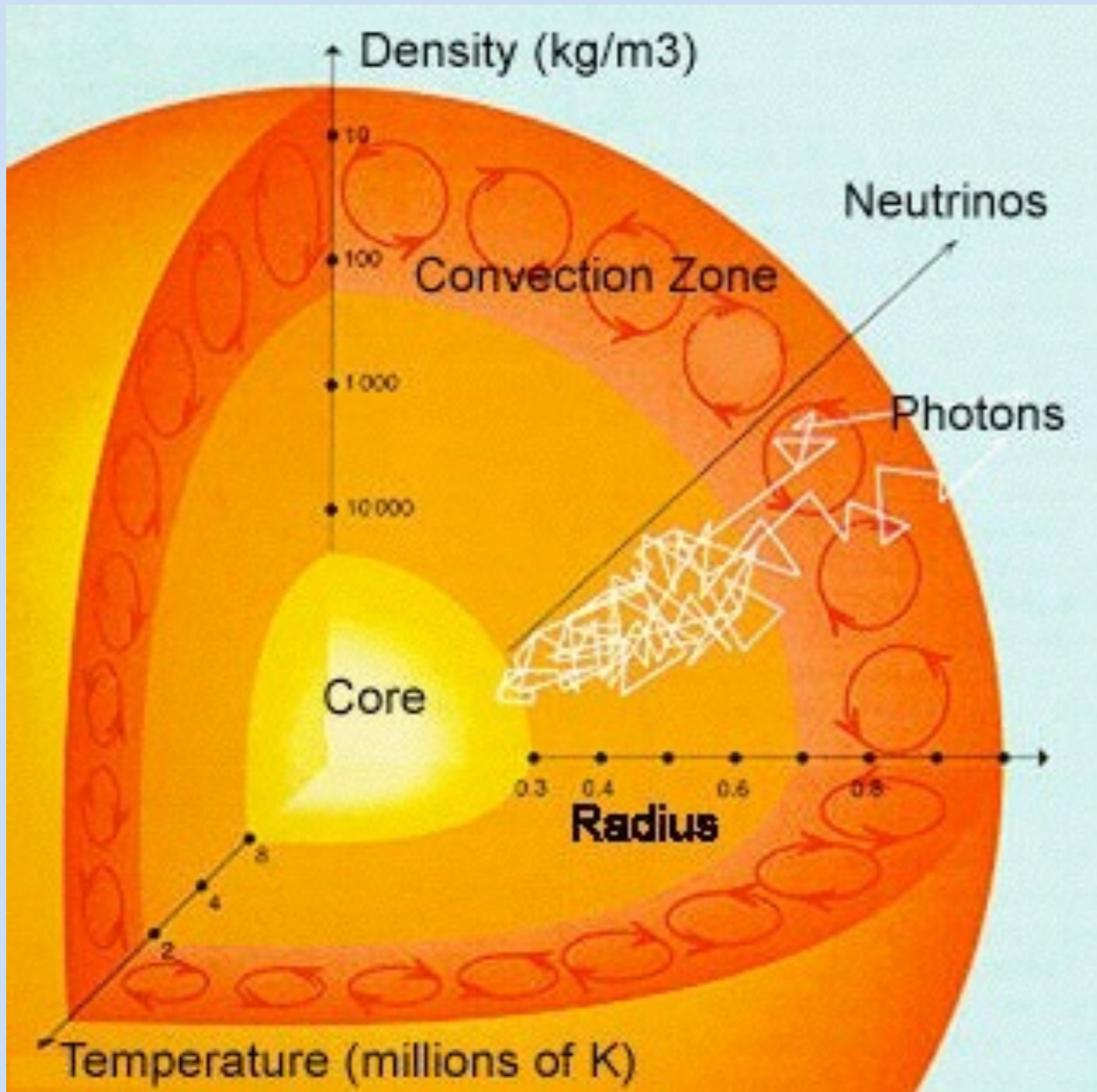
The glue to connect them all.

Helicity in the Sun



Helicity in the Sun

Nonalignment of rotation and gravity

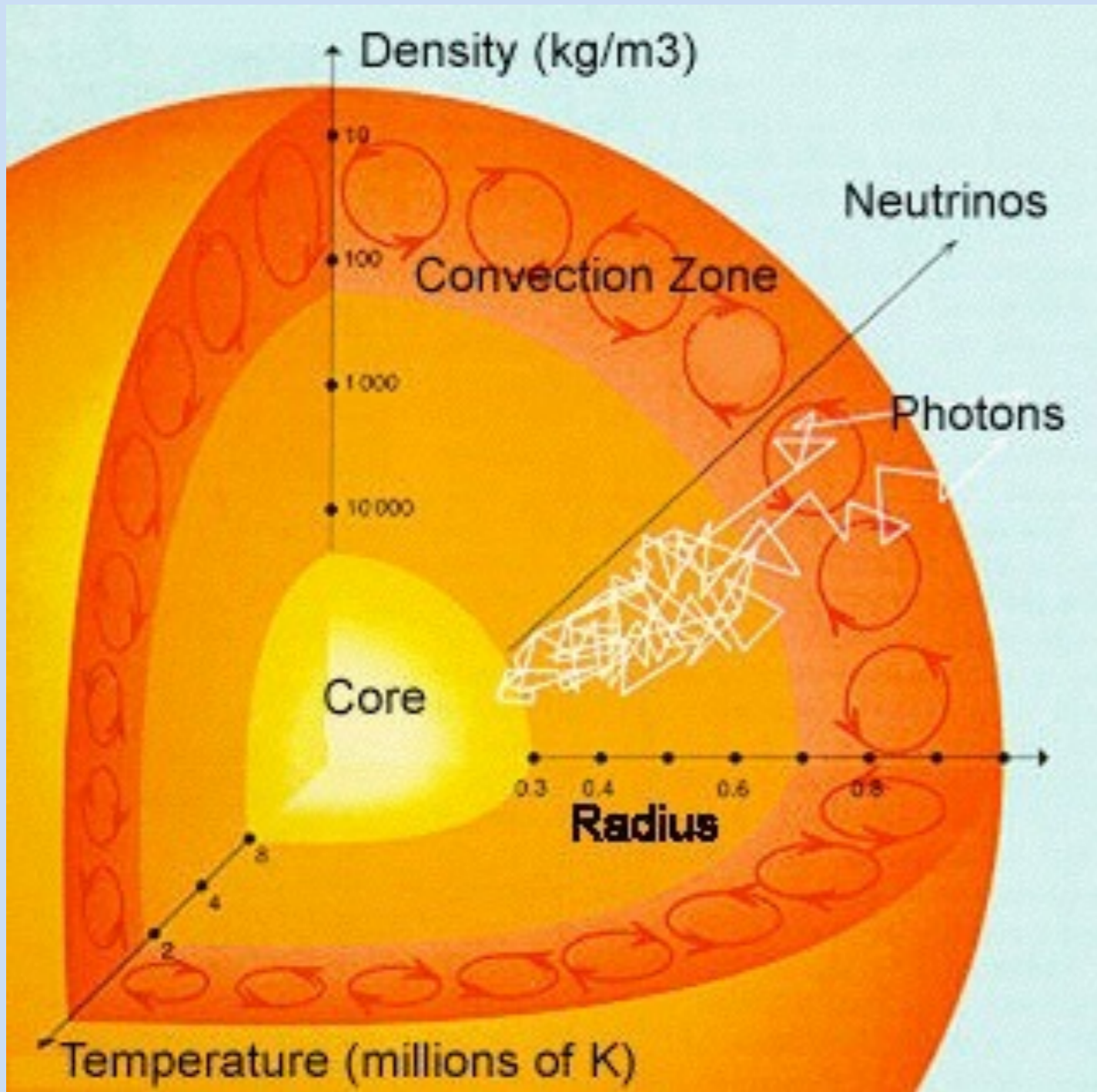


Helicity in the Sun

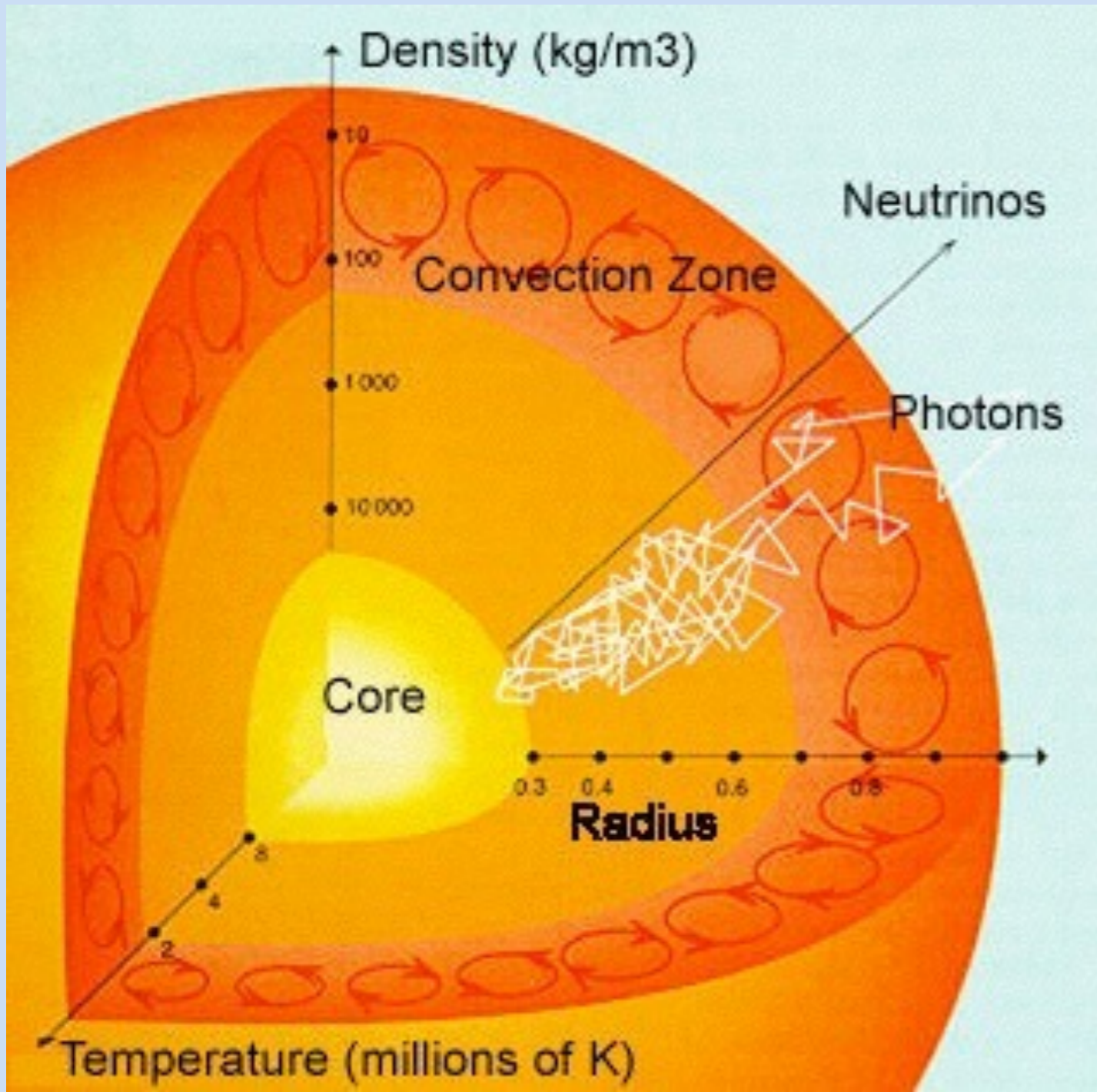
Nonalignment of rotation and gravity



Kinetic helicity



Helicity in the Sun



Nonalignment of
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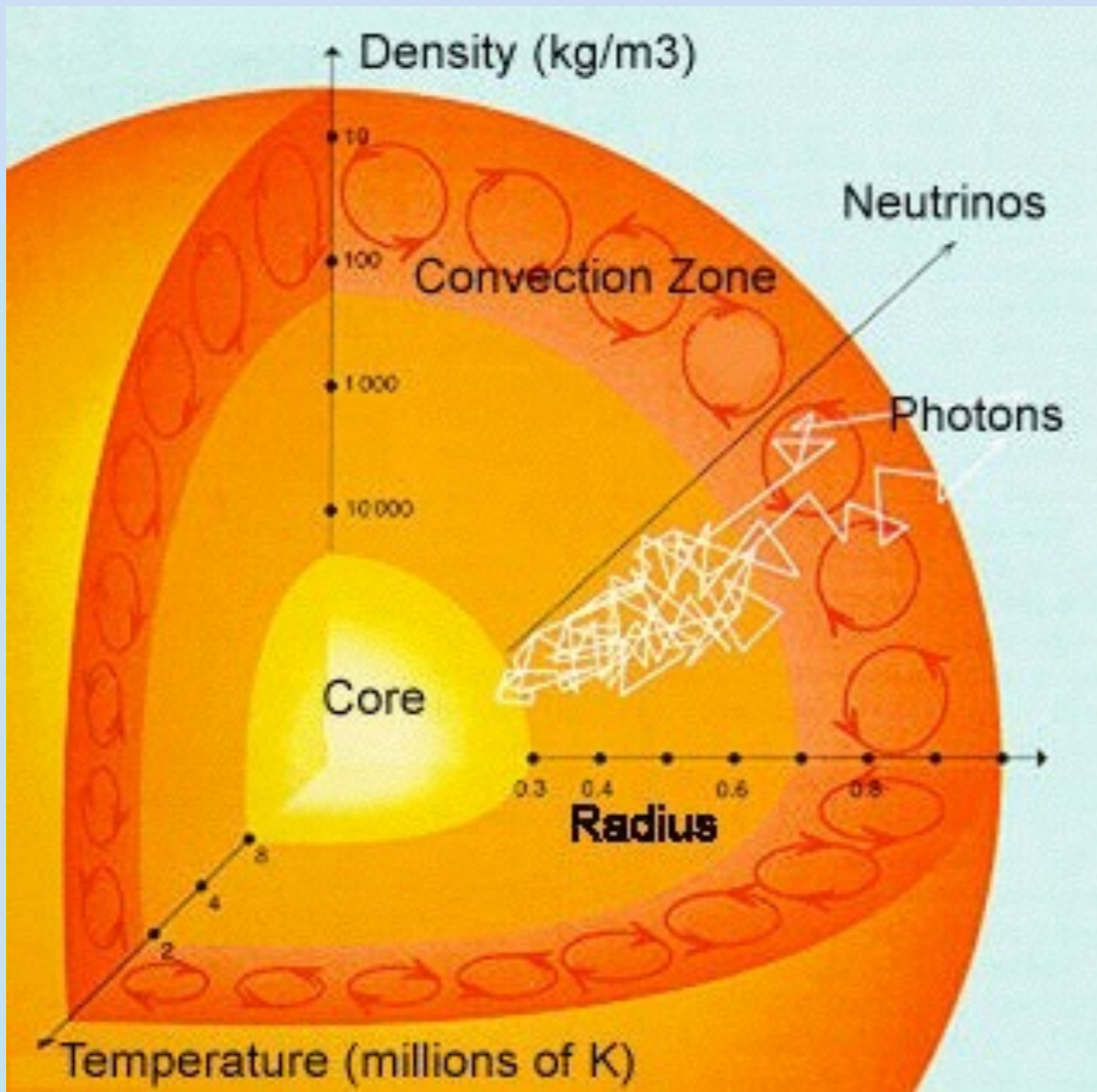


Kinetic helicity



Alpha-effect

Helicity in the Sun



Nonalignment of
rotation and gravity



Kinetic helicity

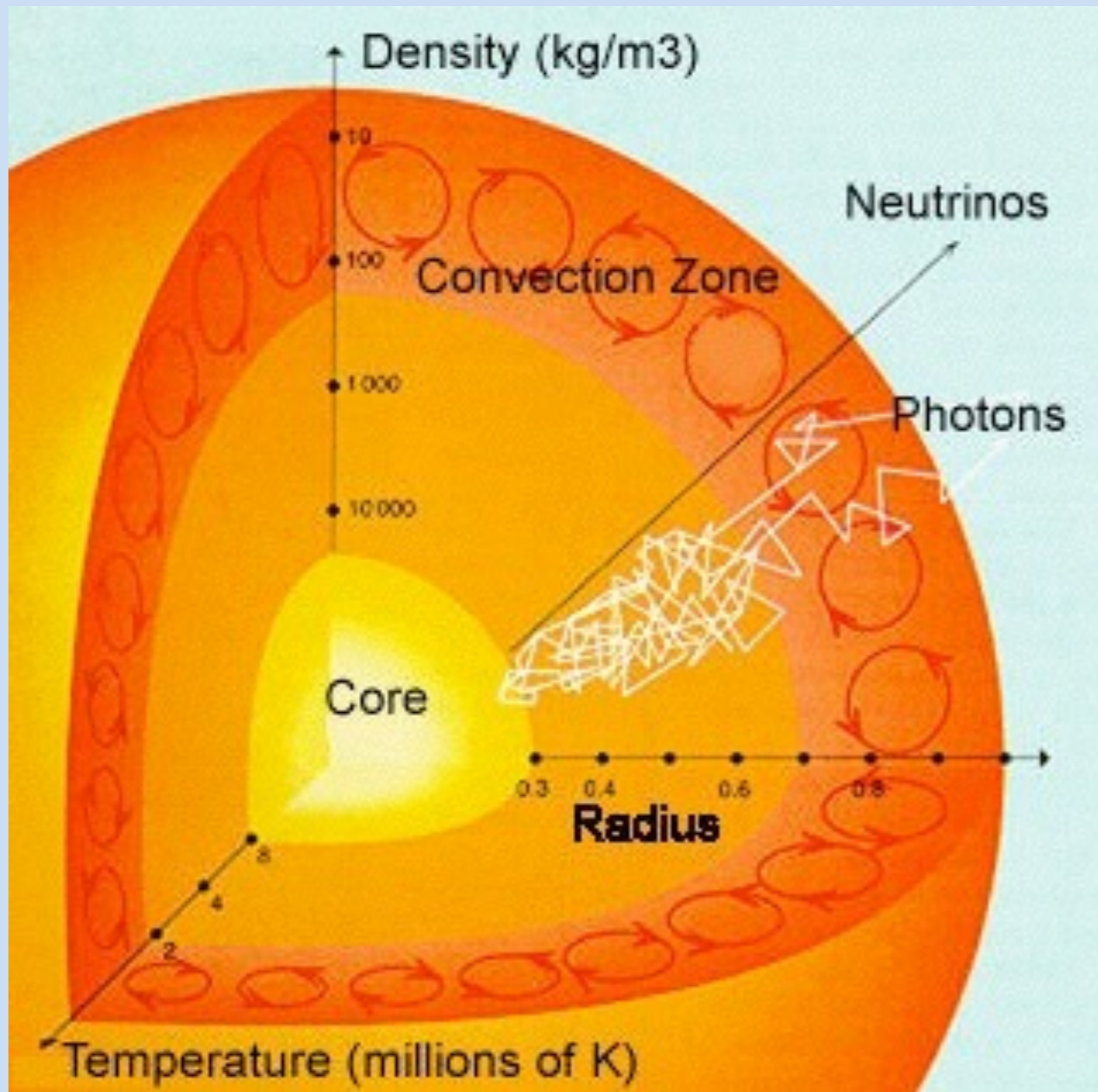


Alpha-effect



Magnetic helicity
+ catastroph. quenching

Helicity in the Sun



Nonalignment of
rotation and gravity



Kinetic helicity



Alpha-effect



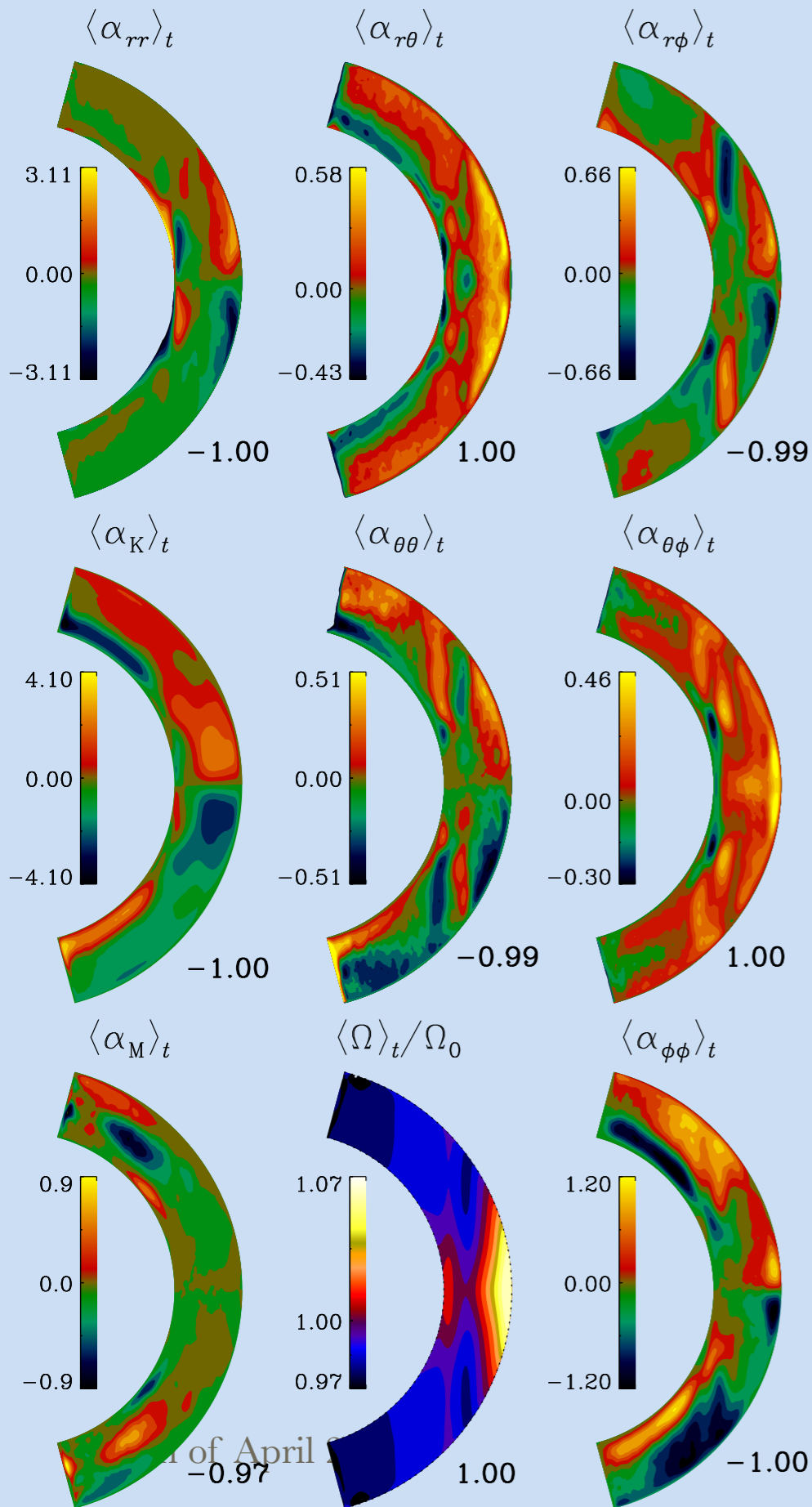
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Space weather

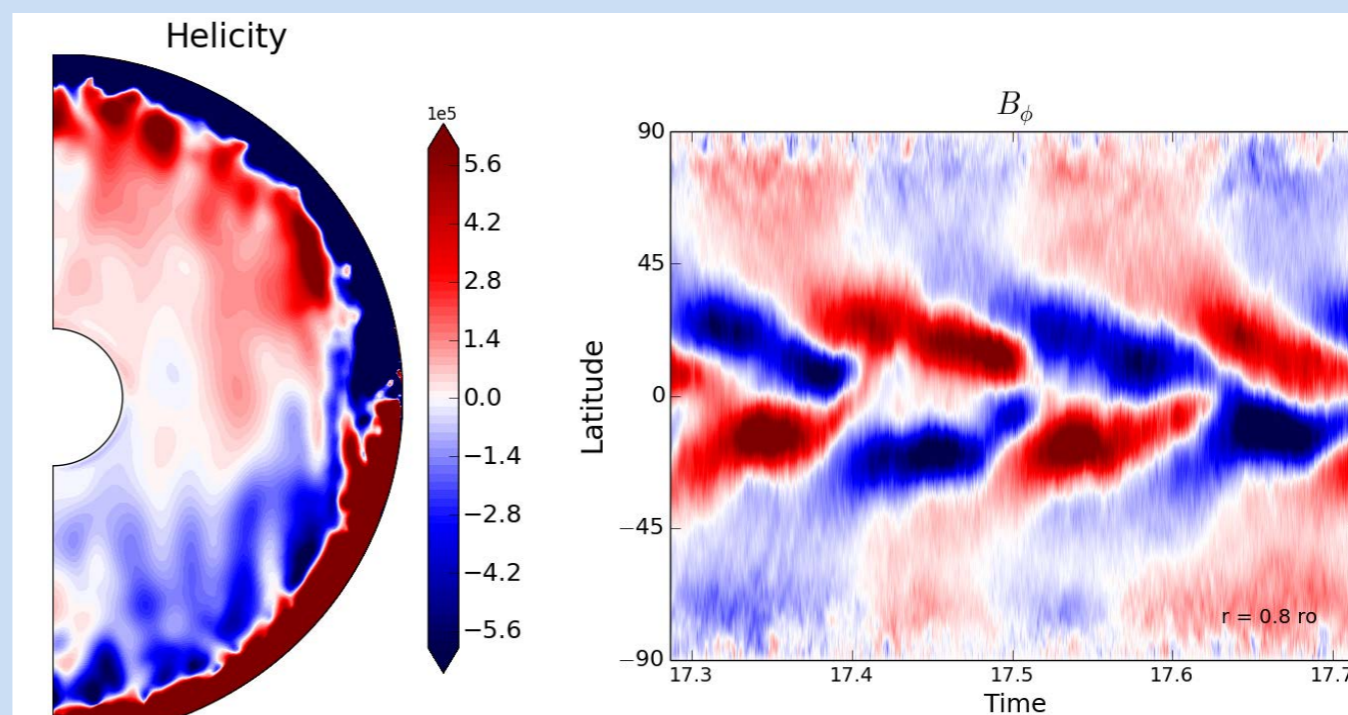
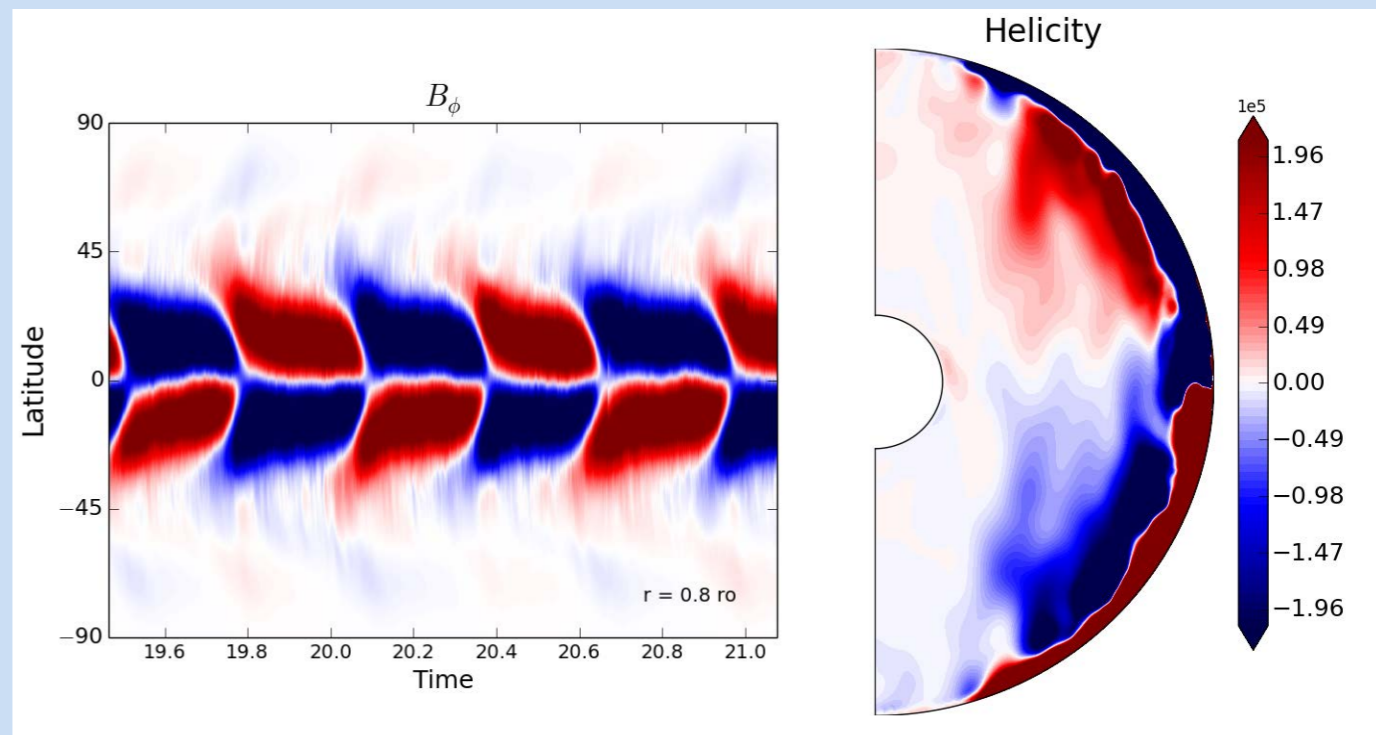
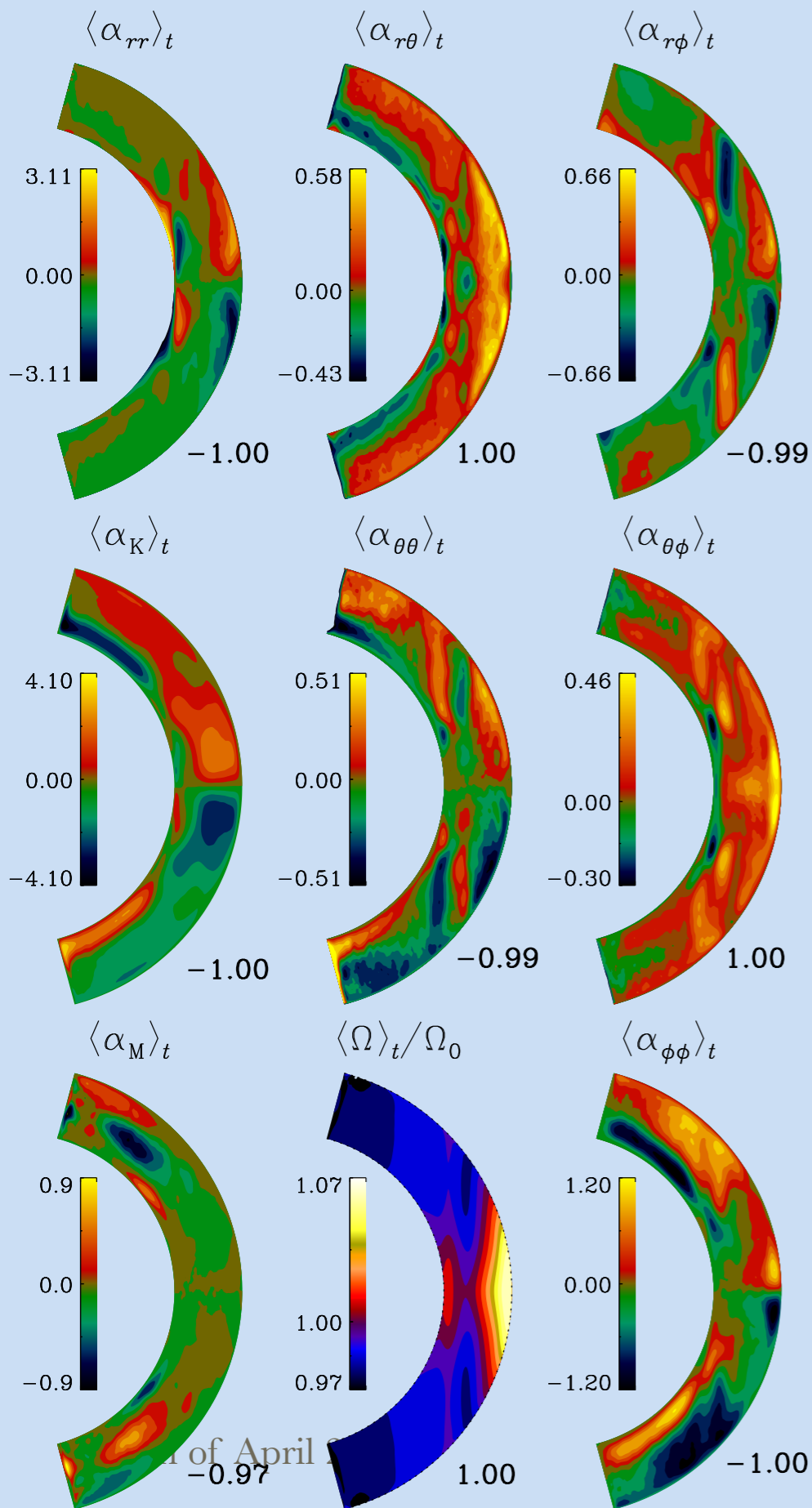
Alpha effect

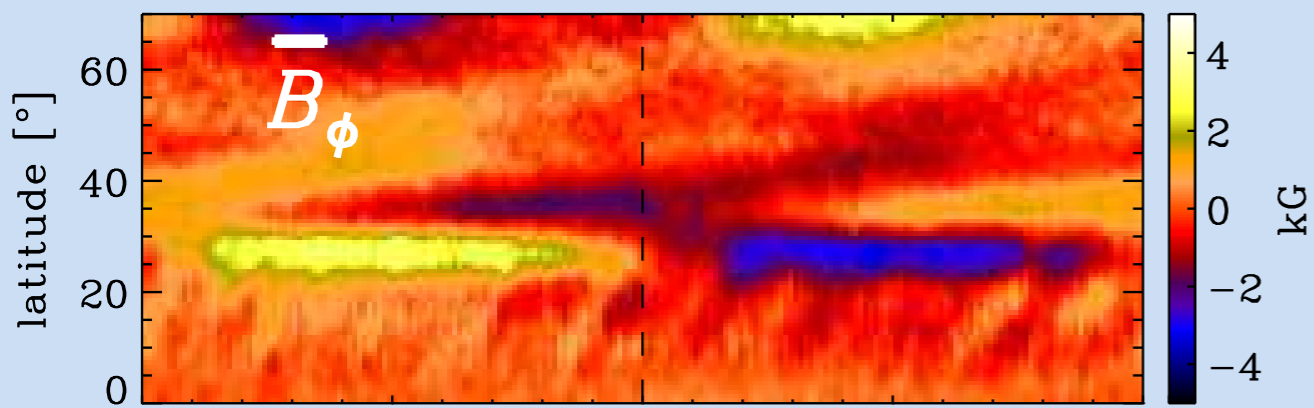
$$\alpha = \frac{\tau_c}{3} \left(-\overline{\boldsymbol{\omega} \cdot \mathbf{u}} + \frac{\overline{\mathbf{j} \cdot \mathbf{b}}}{\bar{\rho}} \right)$$



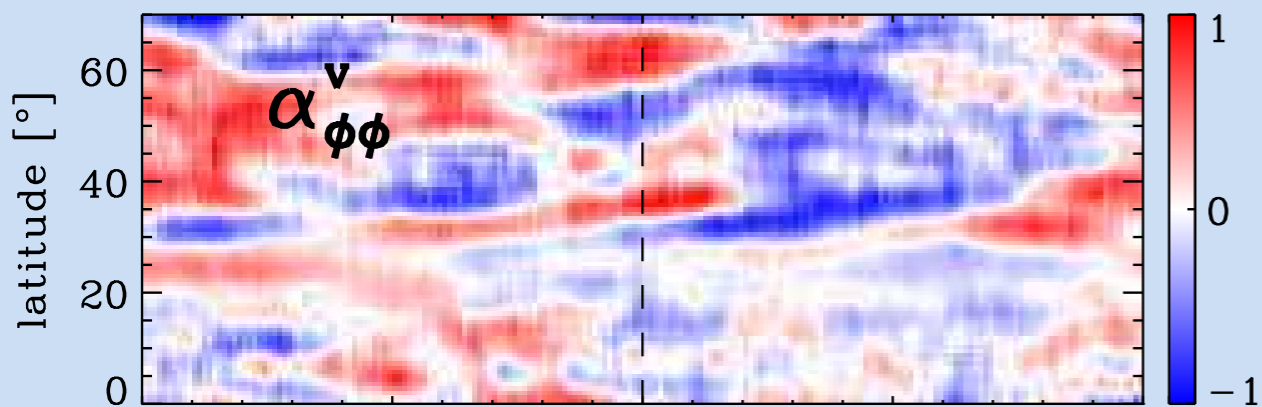
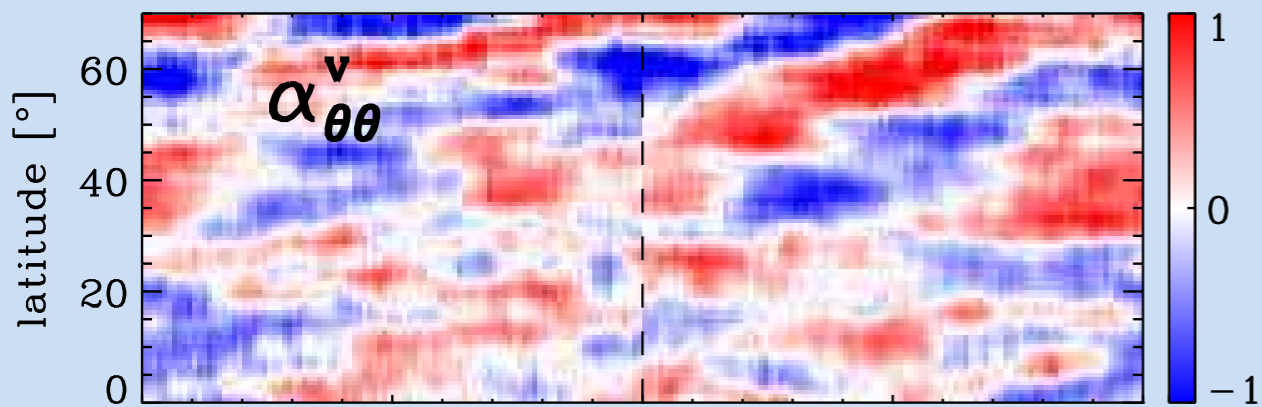
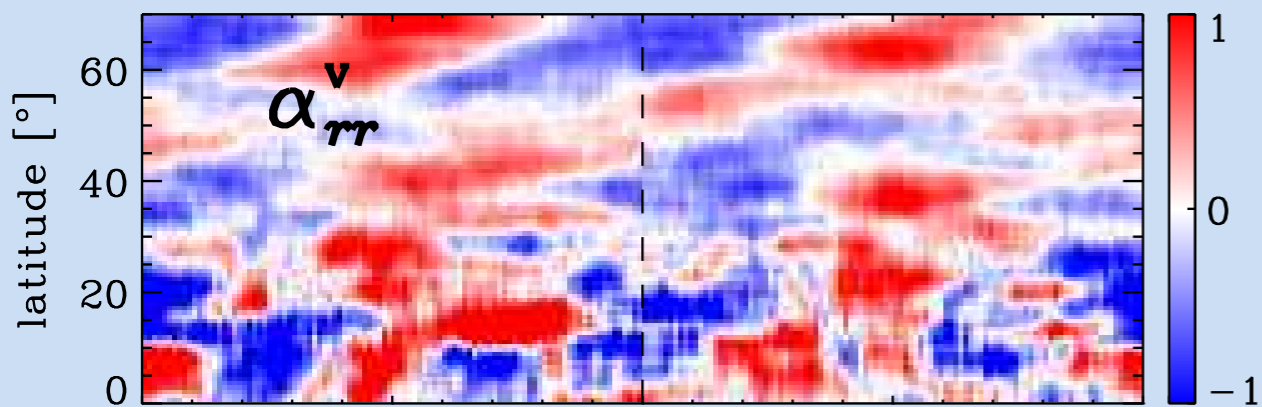
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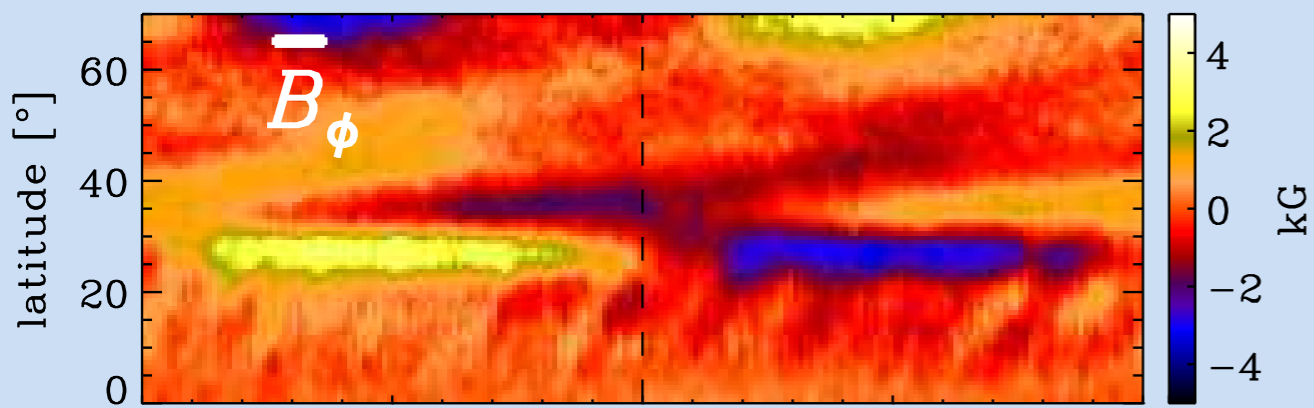
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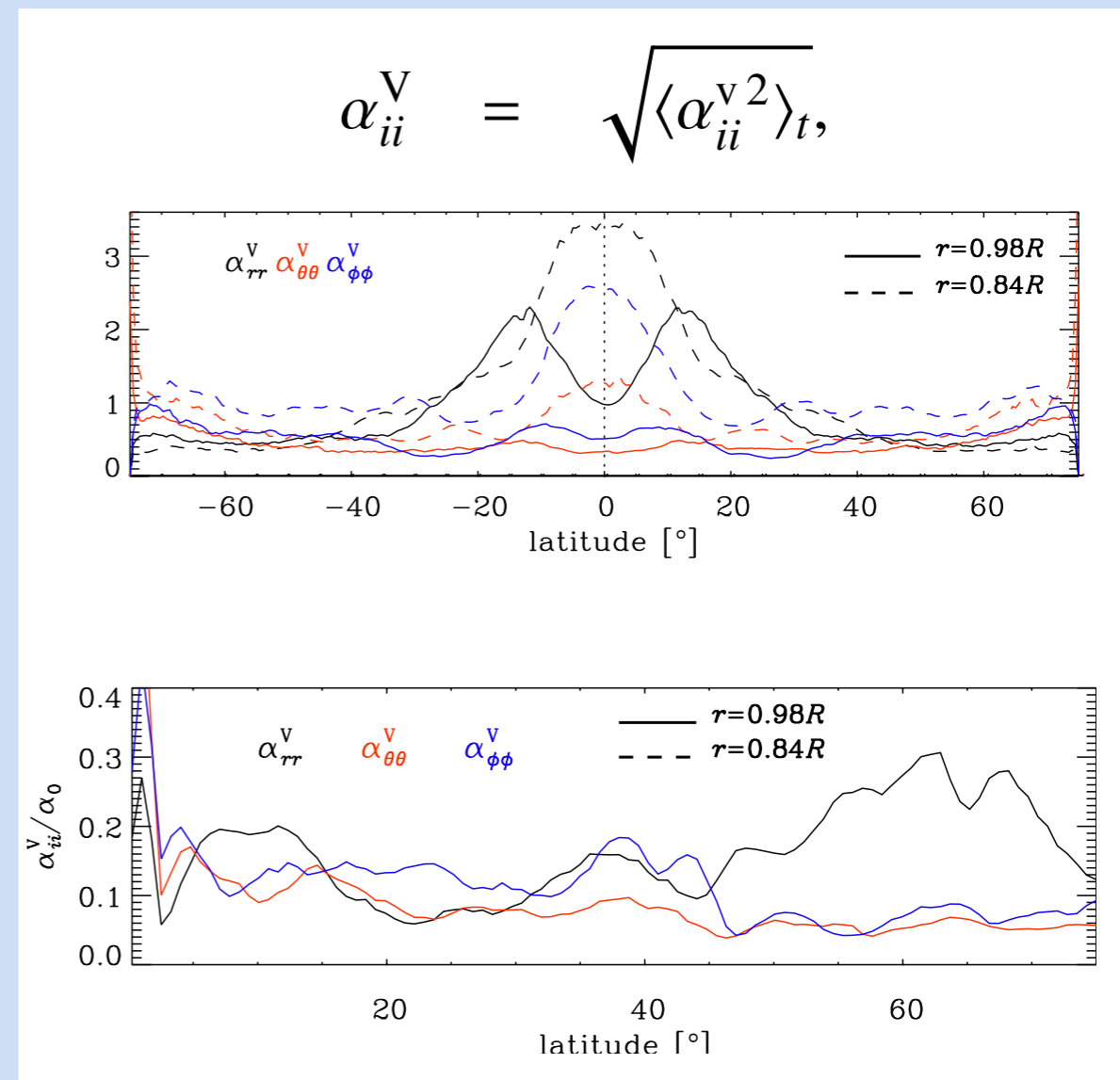
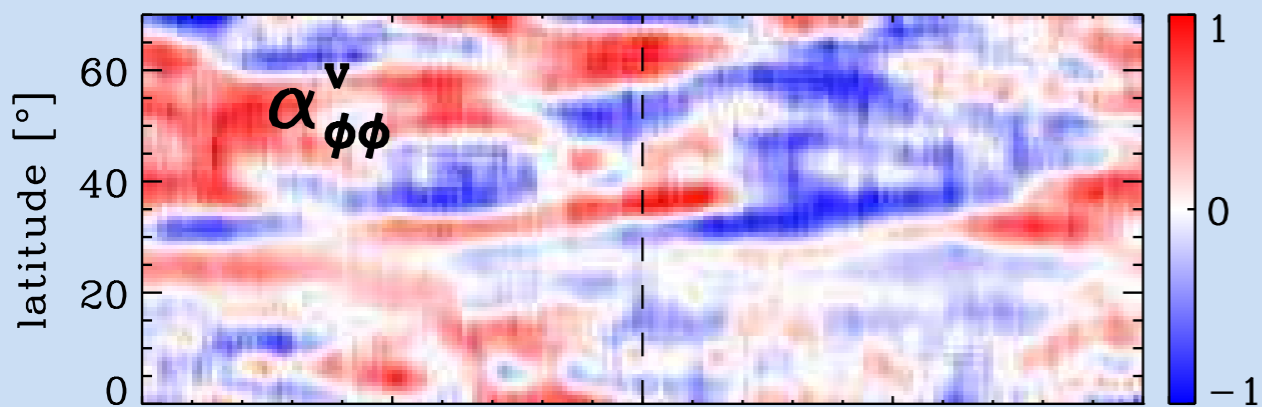
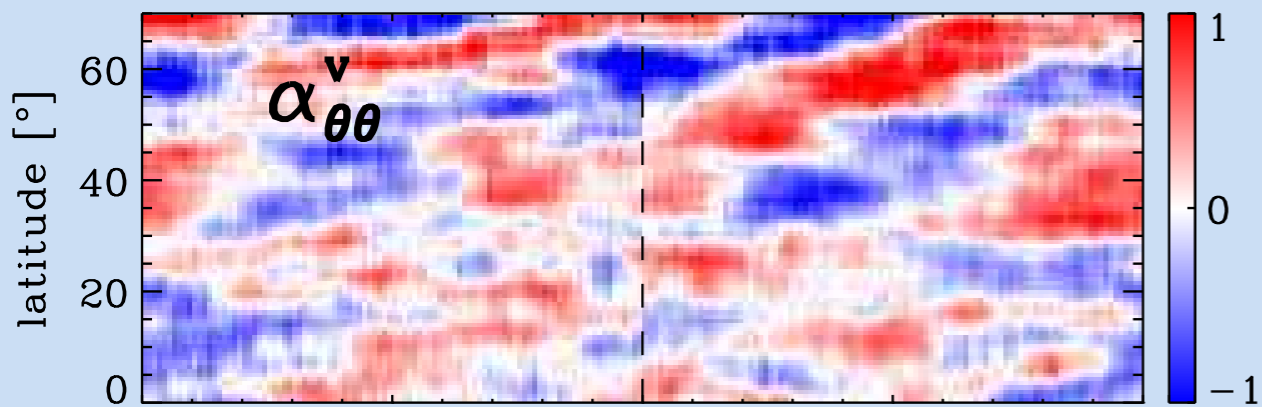
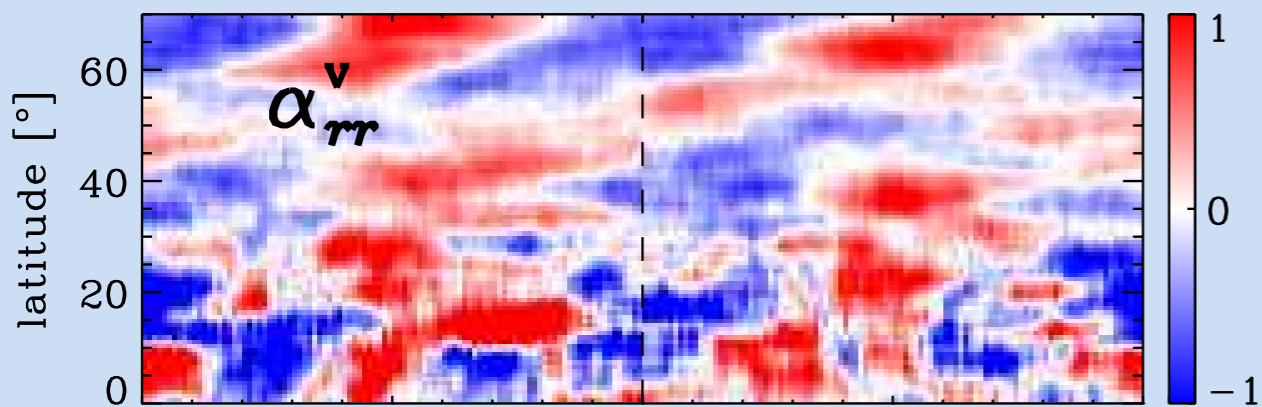


$$\alpha = \langle \alpha \rangle_t + \alpha^v.$$

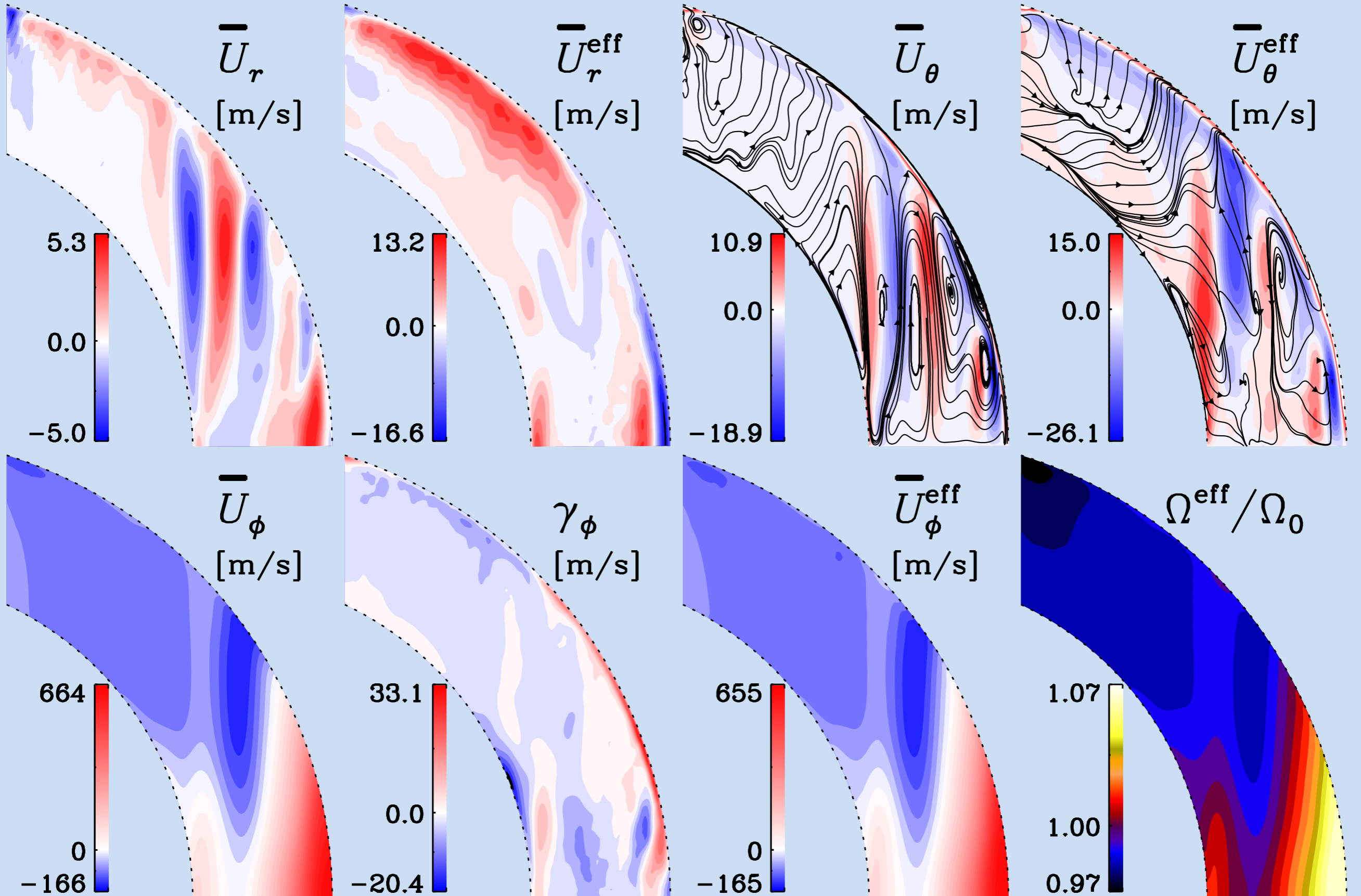




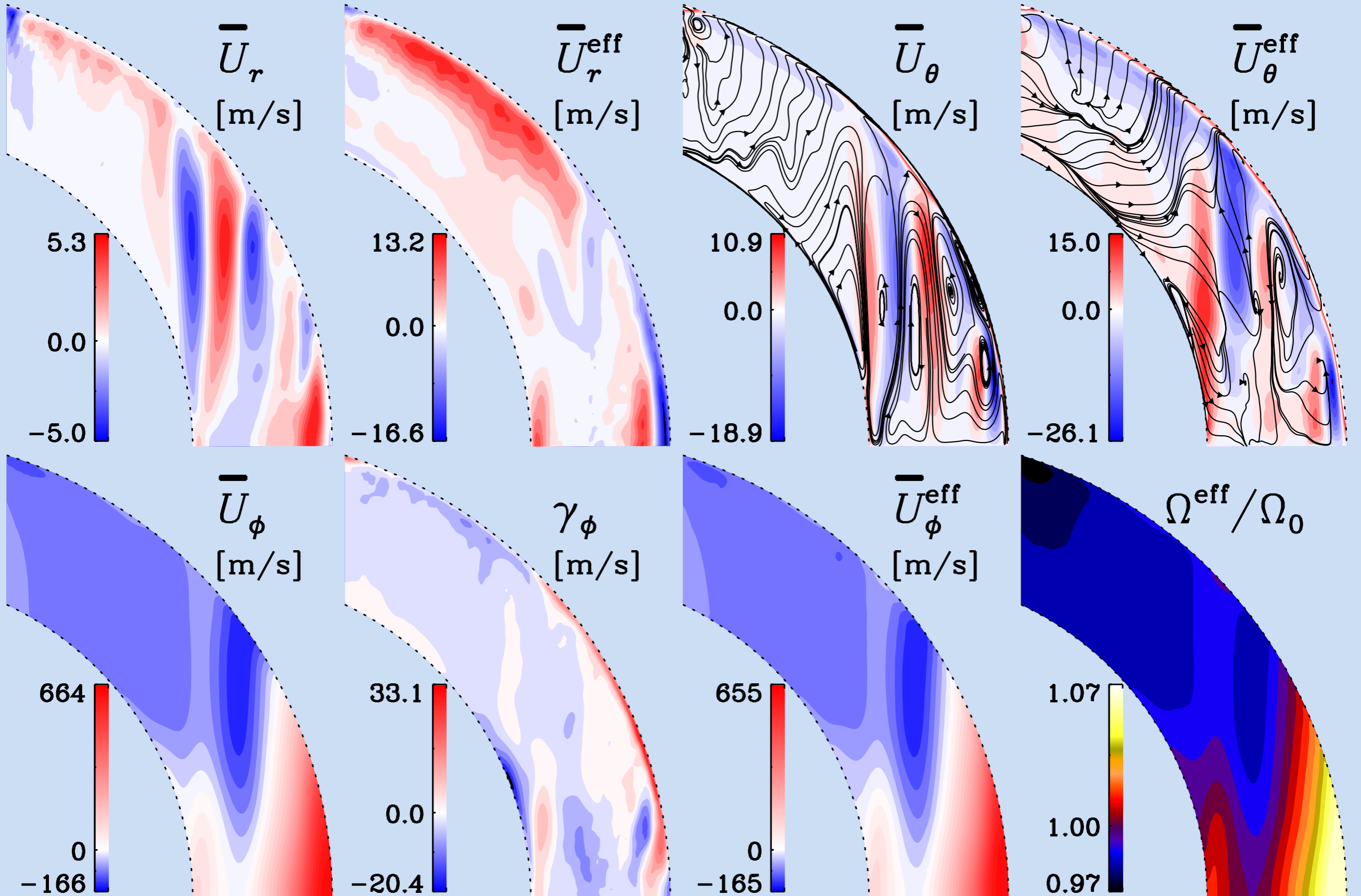
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Turbulent pumping



Turbulent pumping



4th of April 2016

Space Climate 6, Levi, Fin

Warnecke et al., 2016

Dynamo is quenched at high magnetic Reynolds numbers

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Magnetic helicity fluxes can prevent
dynamo action of being quenched.

(Brandenburg et al. 2009, Blackman and Field 2000, Brandenburg and Sandin 2004).

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Coronal mass ejections might be one possibility
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A realistic boundary condition
for the magnetic field is important.

The Two Layer Model

Lower layer:

Convection zone

Dynamo action

➔ Generation of magnetic field

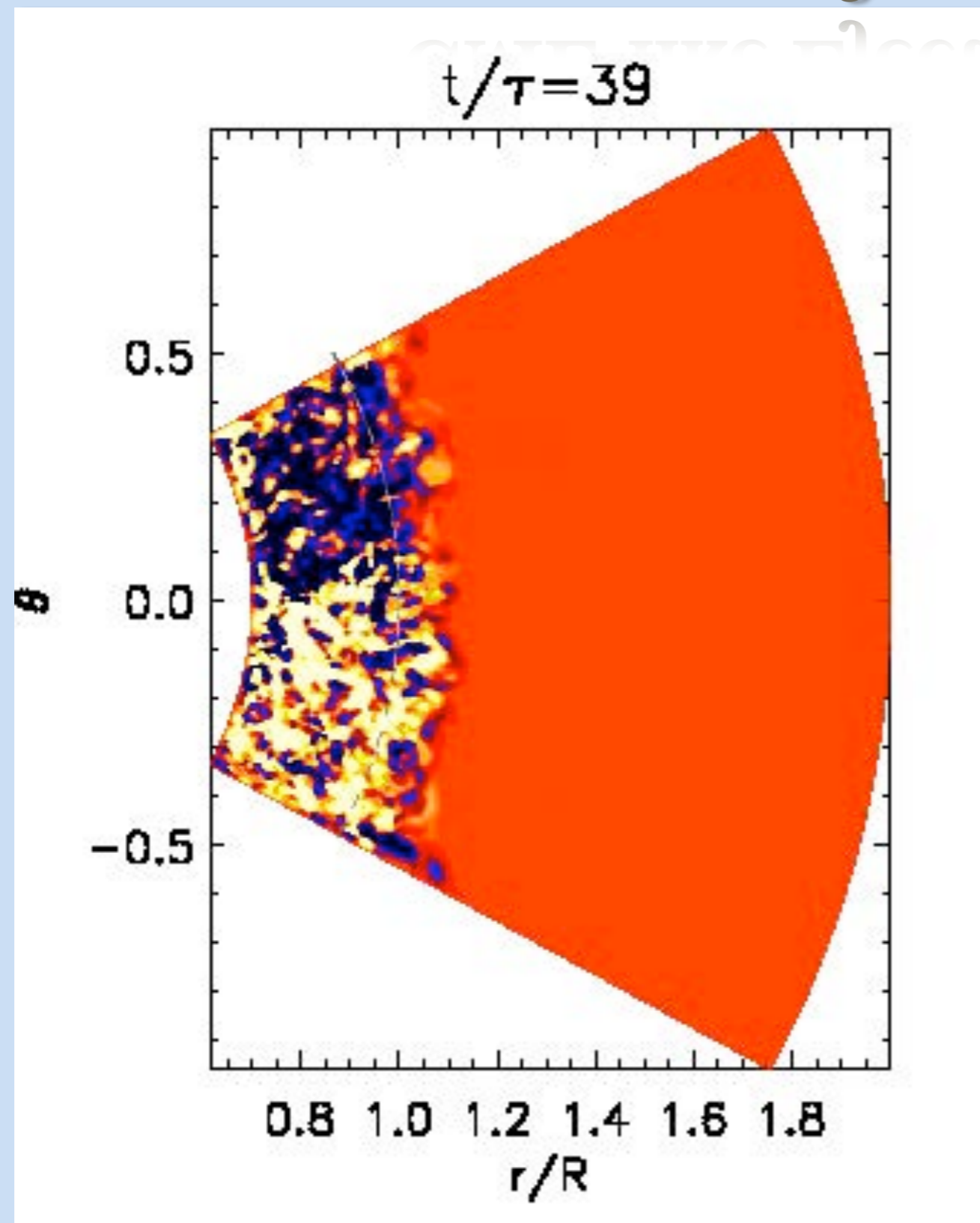
Upper layer:

Simplified coronal model

Magnetic flux emerges from the lower layer and gets ejected.

Both layers are in one simulation.

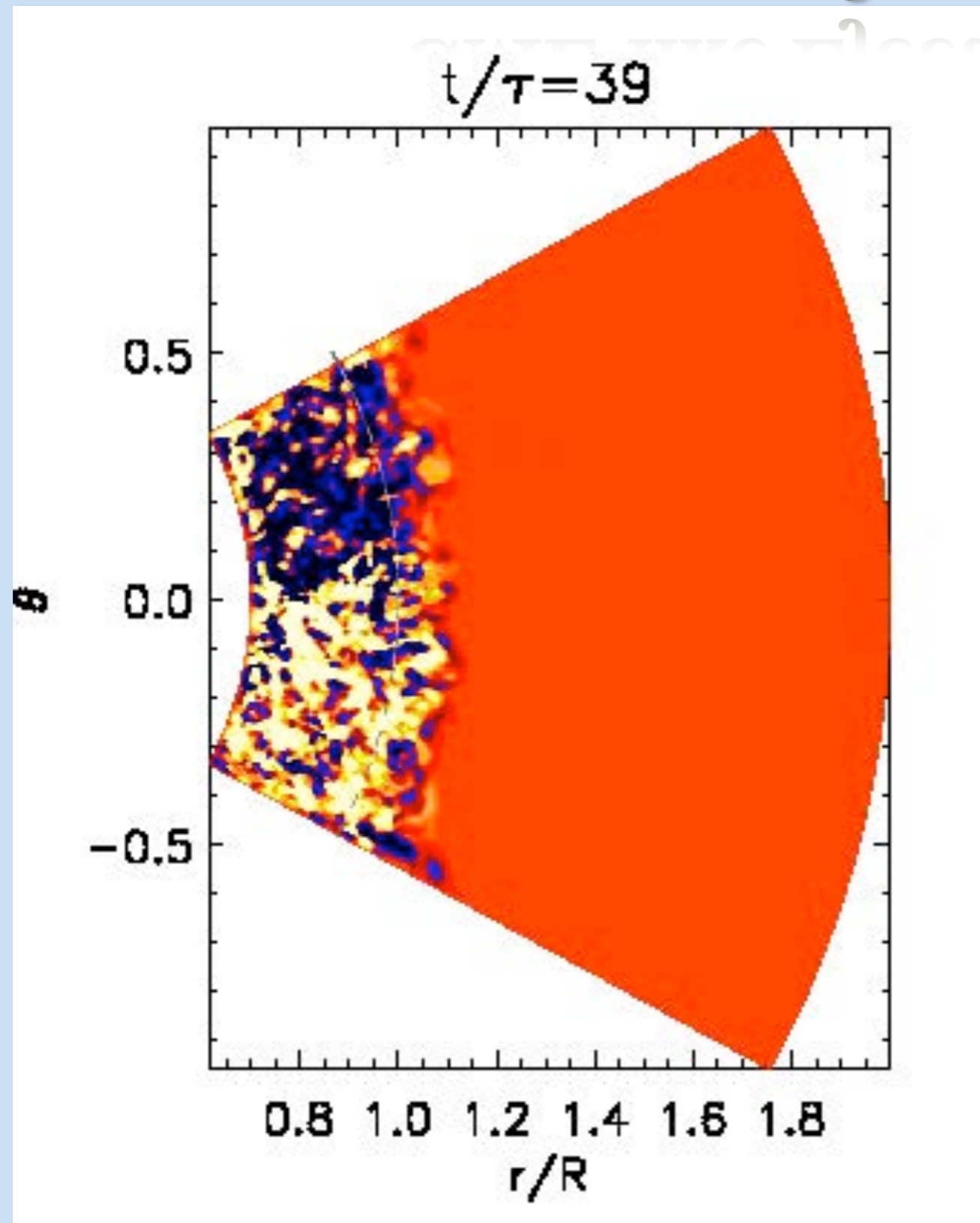
CME-like Ejections



Current helicity
density

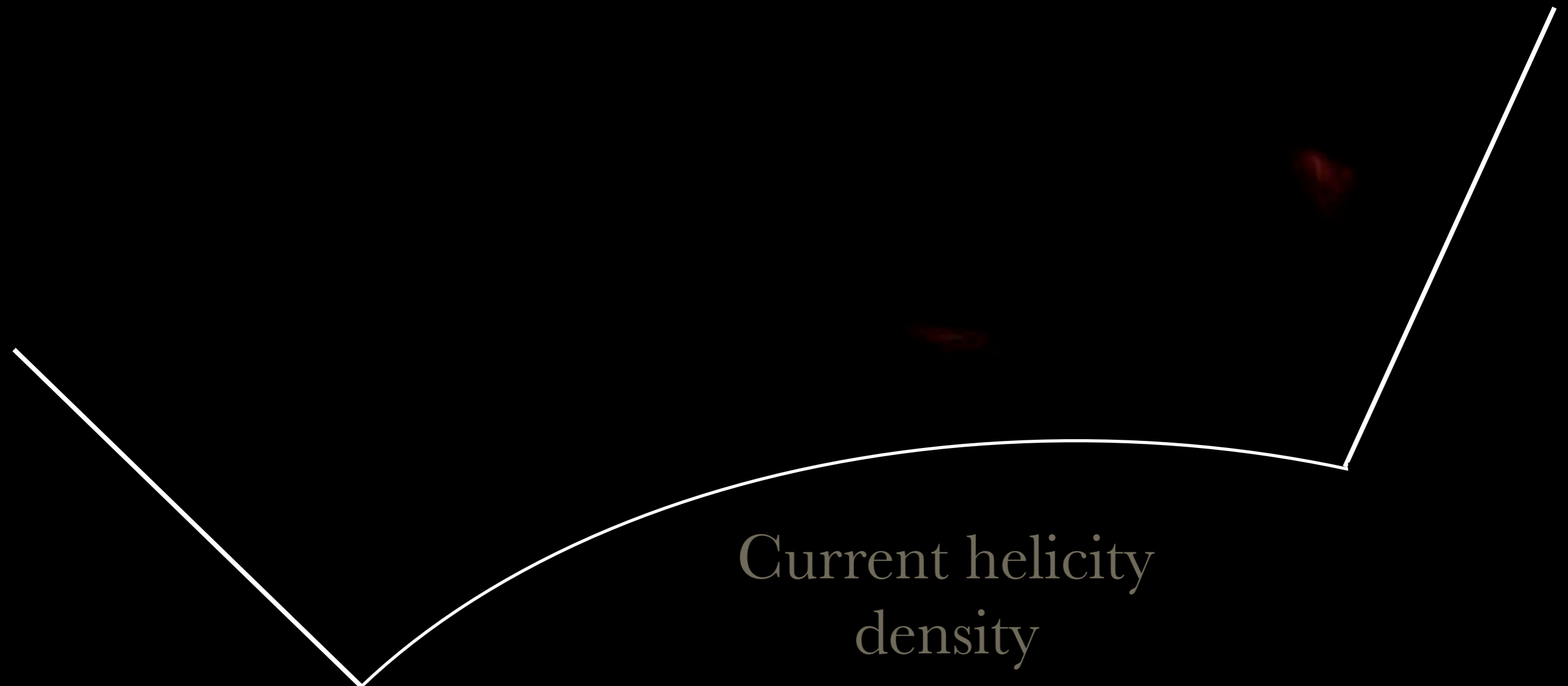
- proxy for magnetic helicity
- gauge invariant
- represent helical structures
- can be measured on the Sun

CME-like Ejections



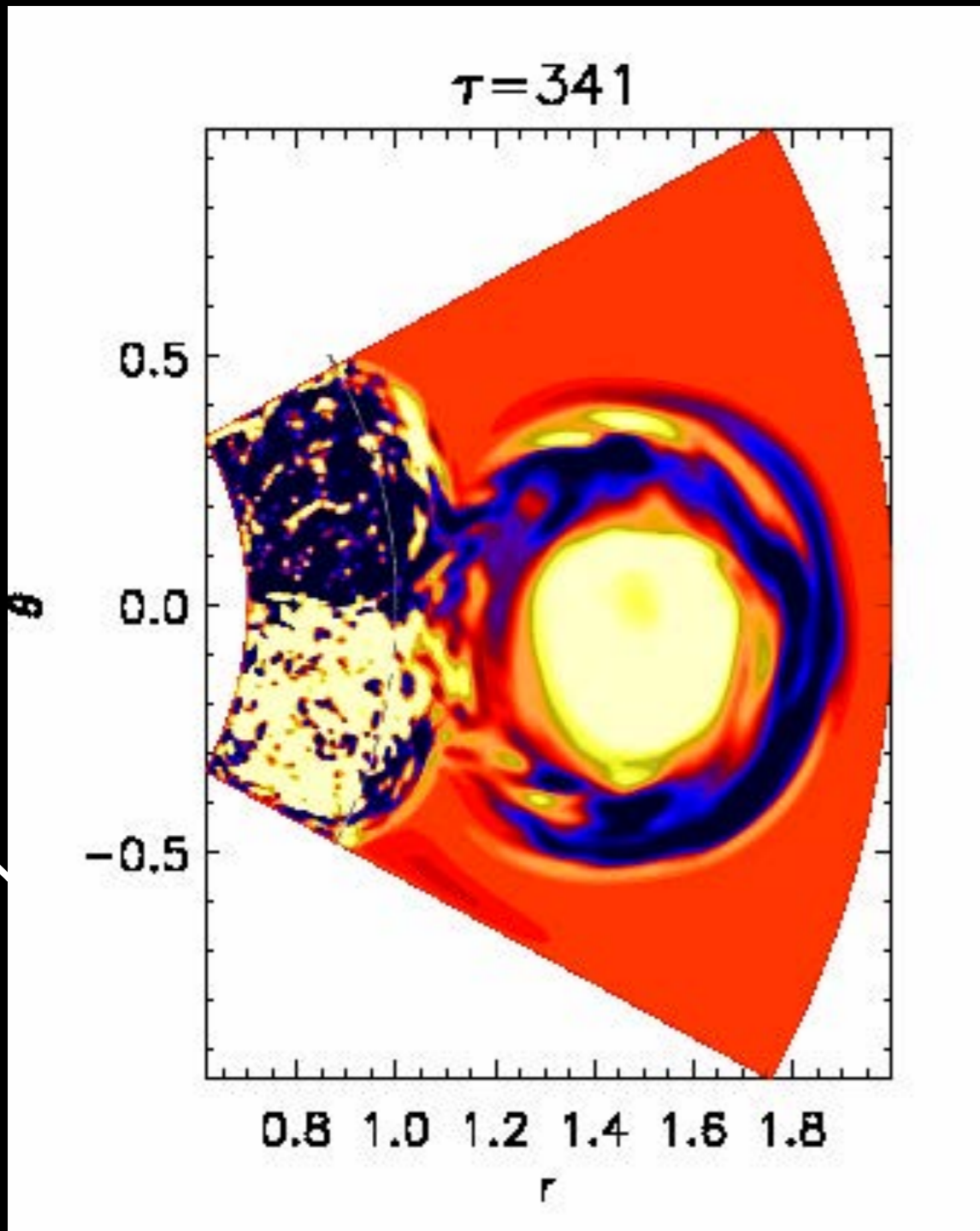
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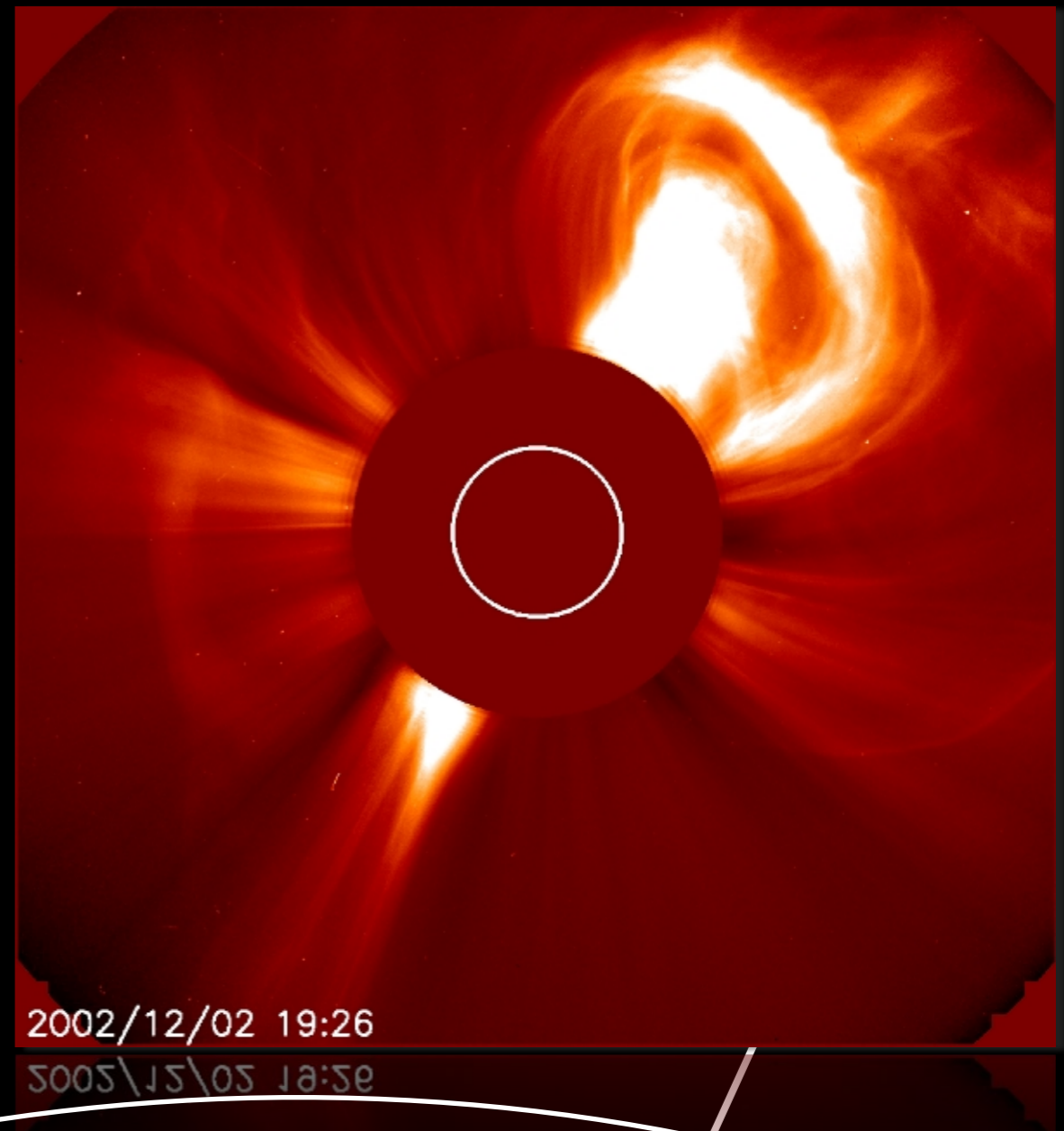
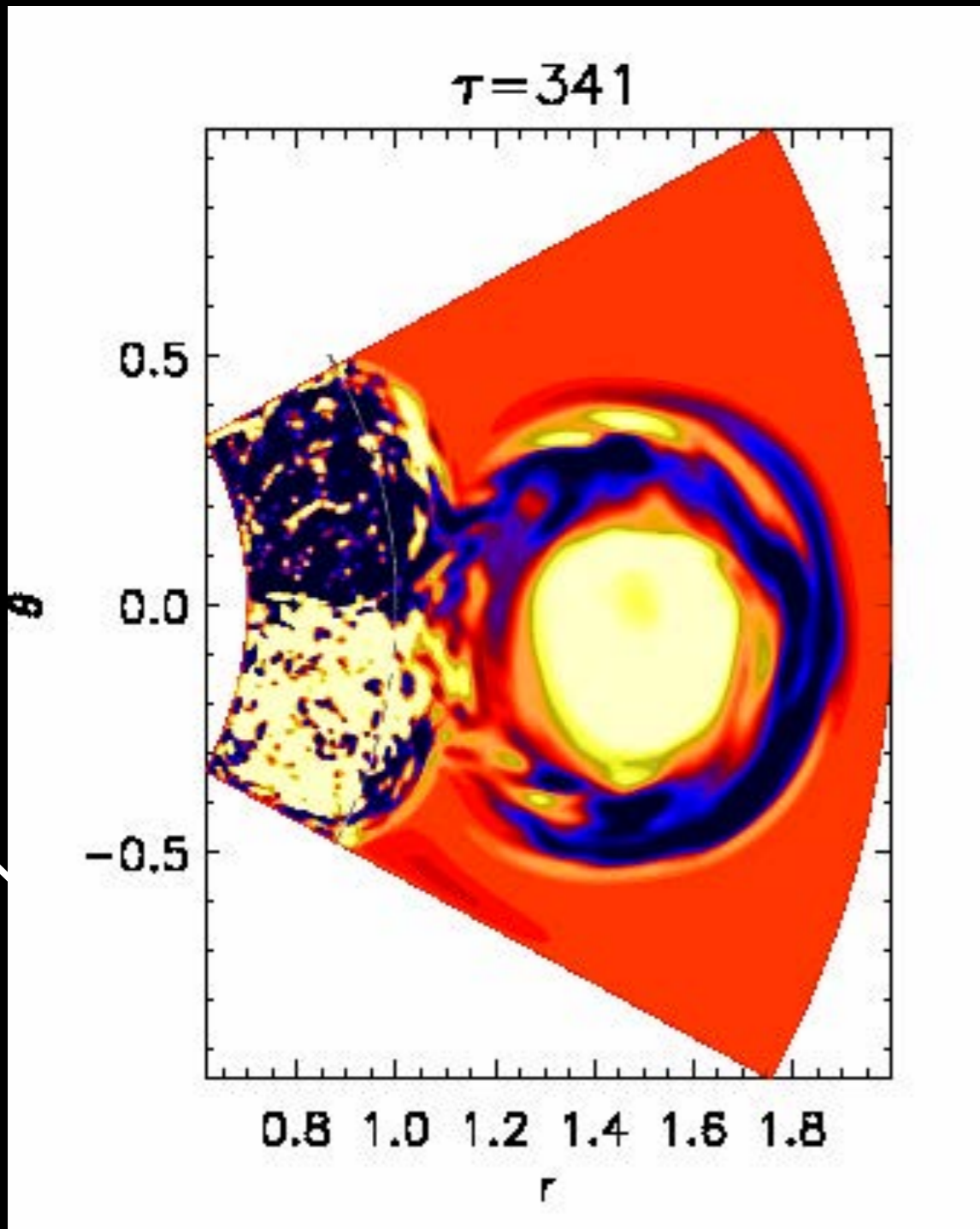
CME-like Ejections



Current helicity
density

0.8 1.0 1.2 1.4 1.6 1.8

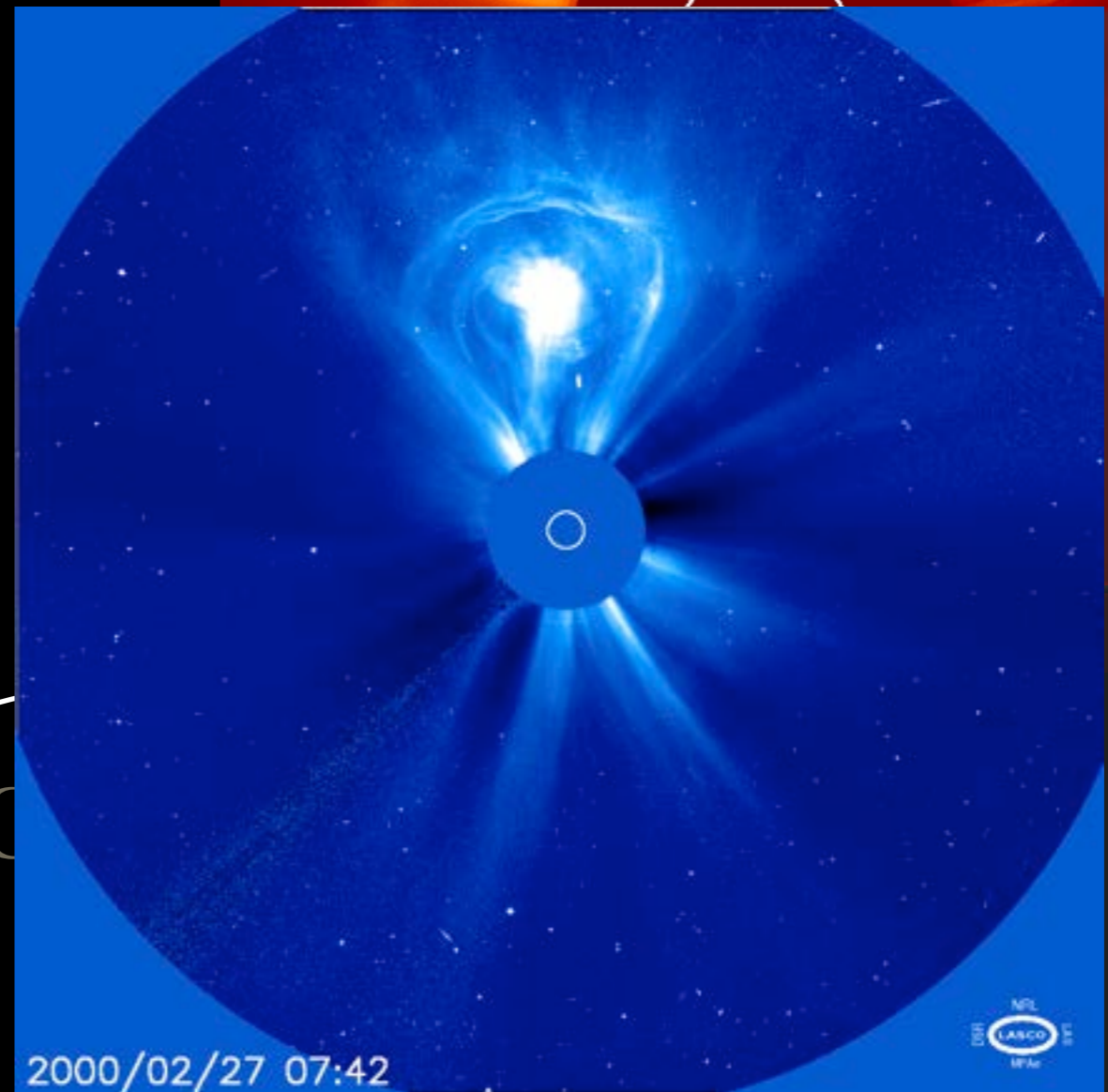
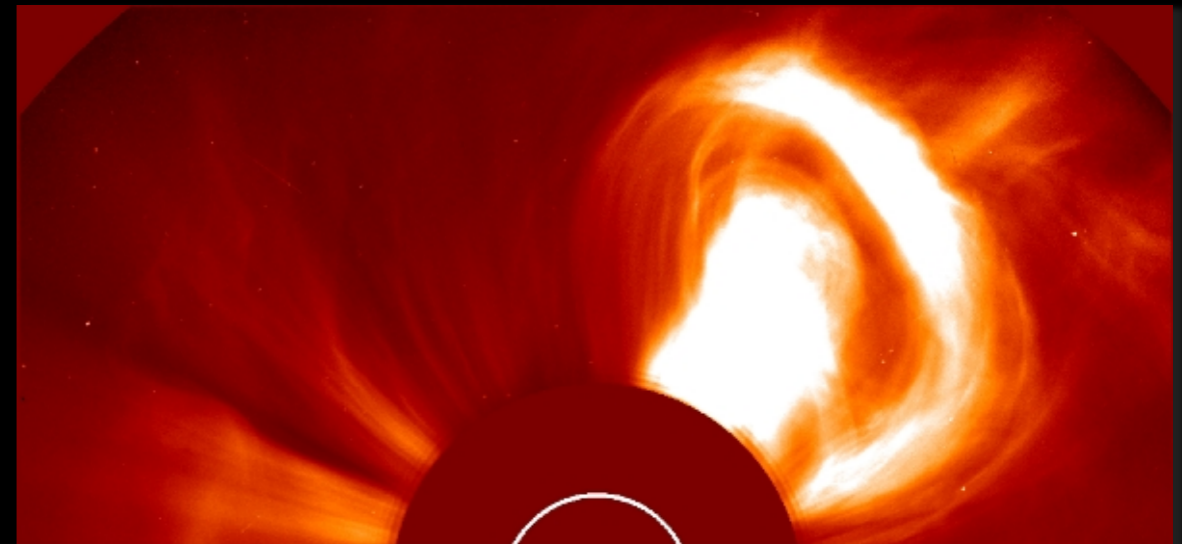
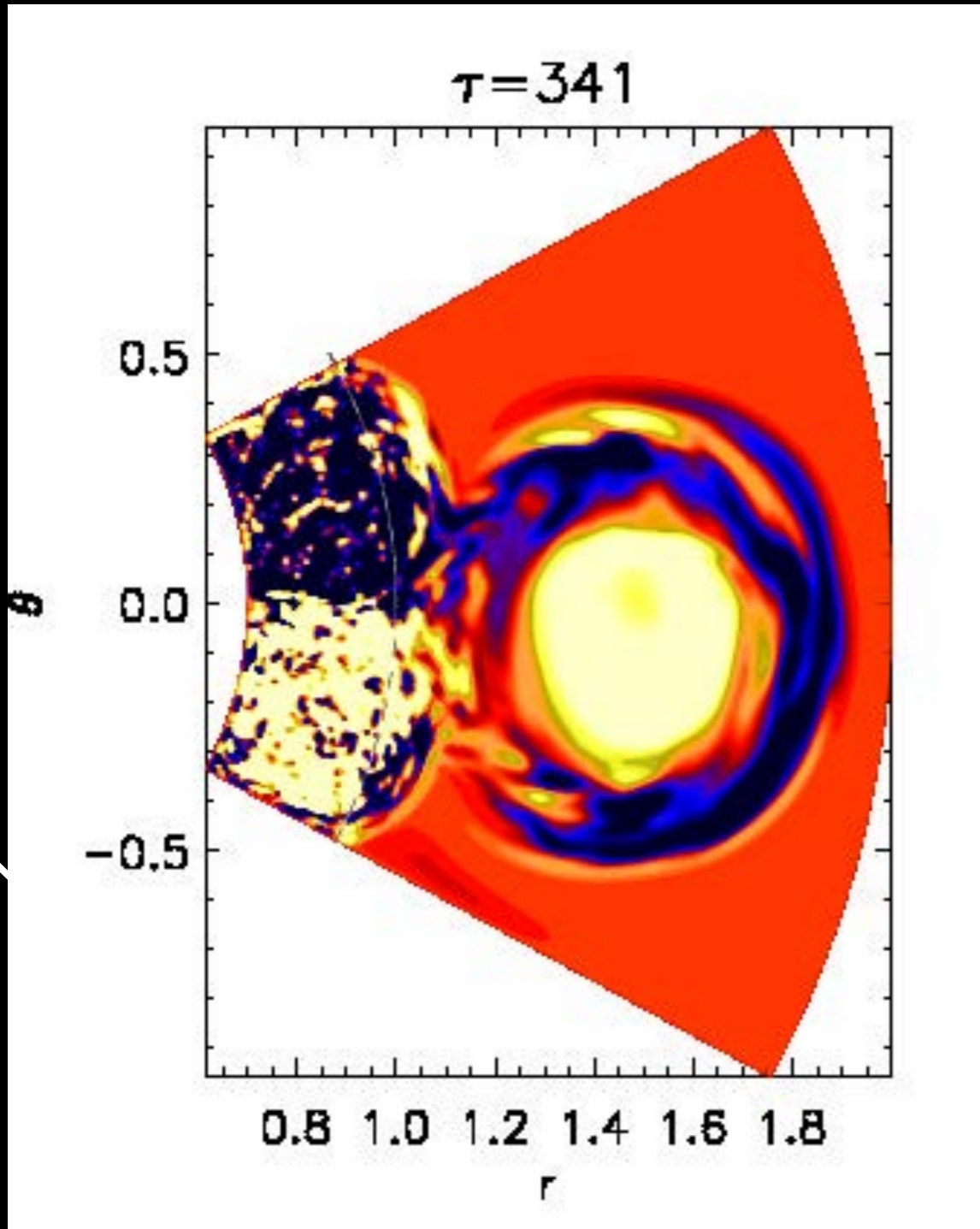
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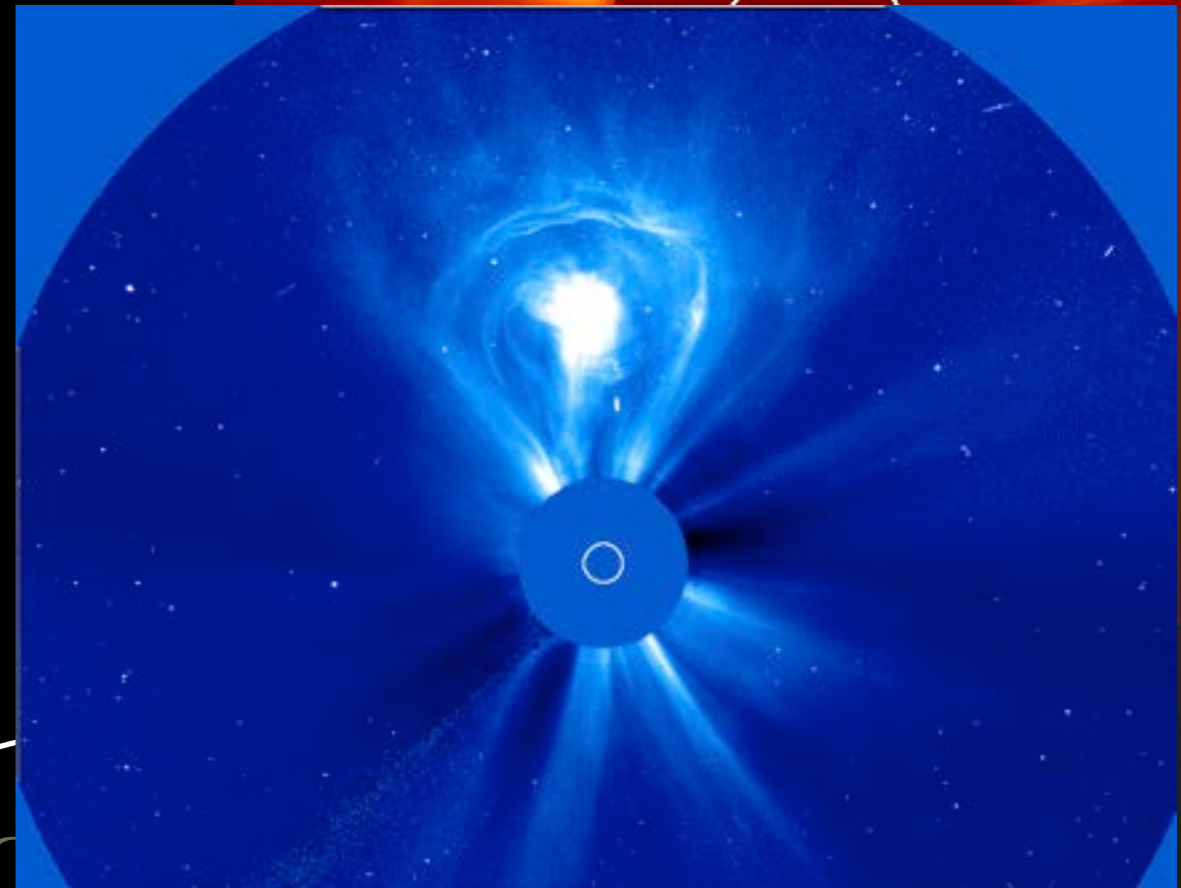
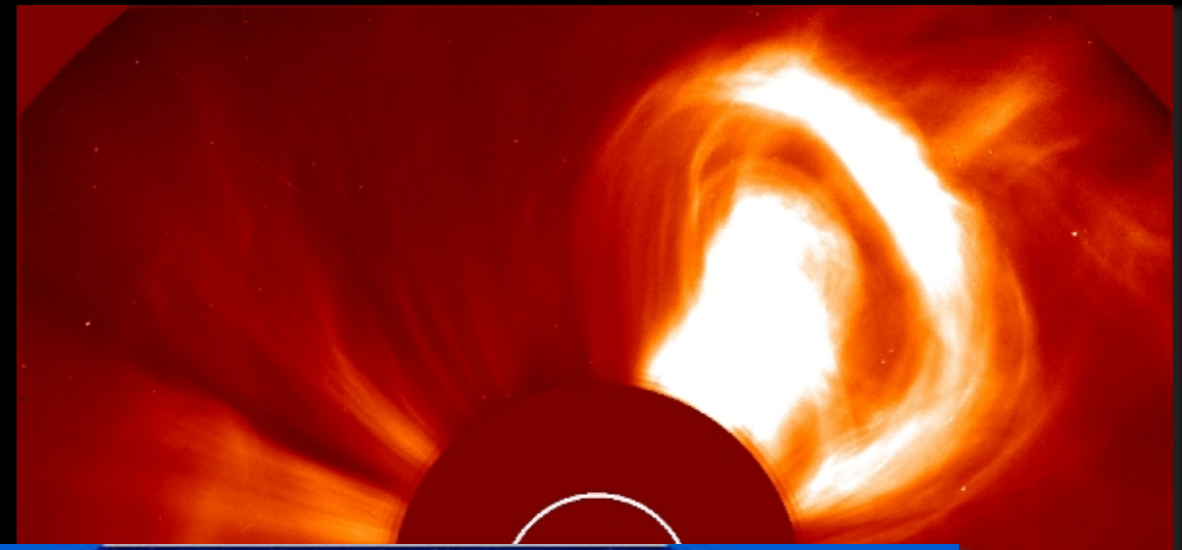
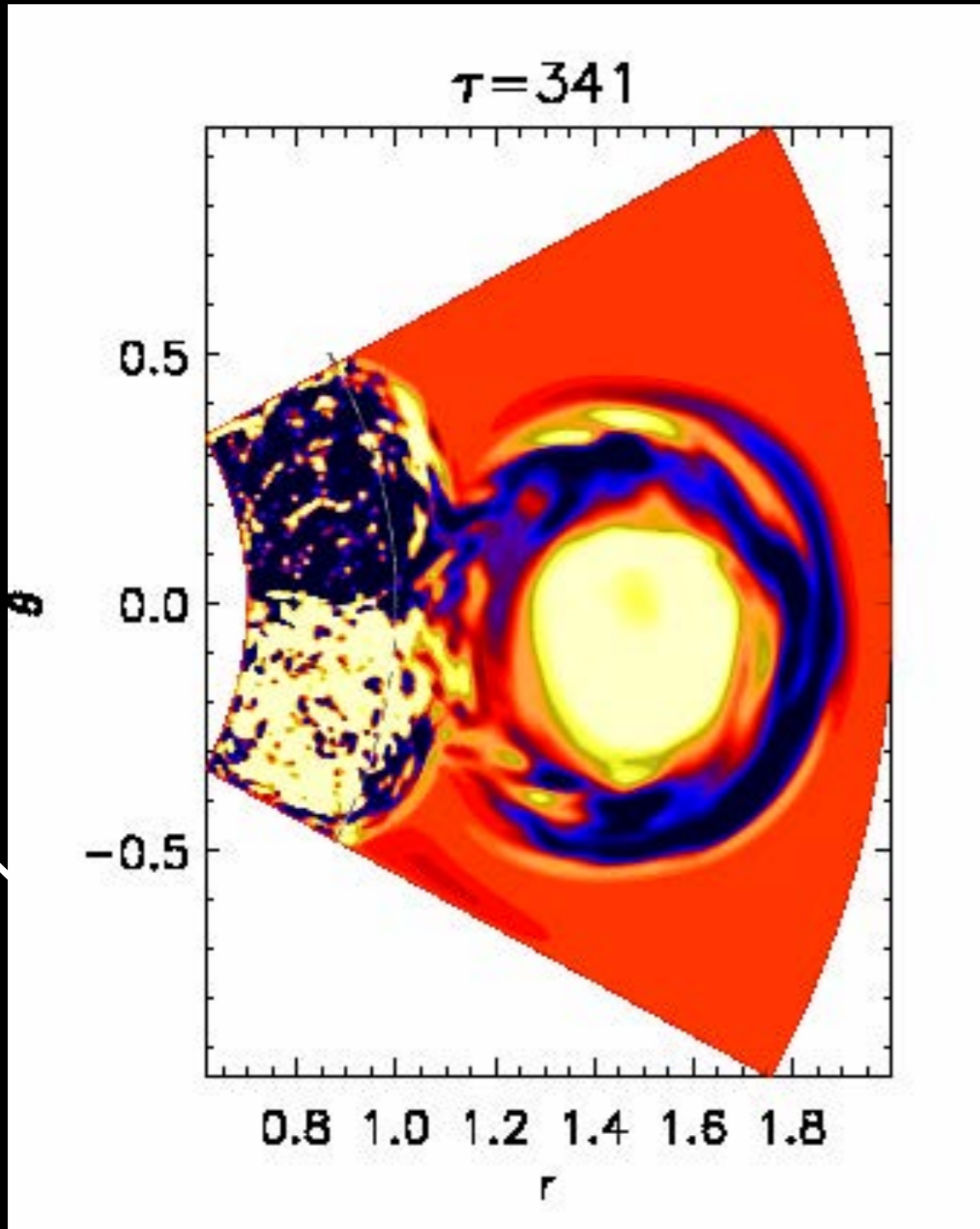
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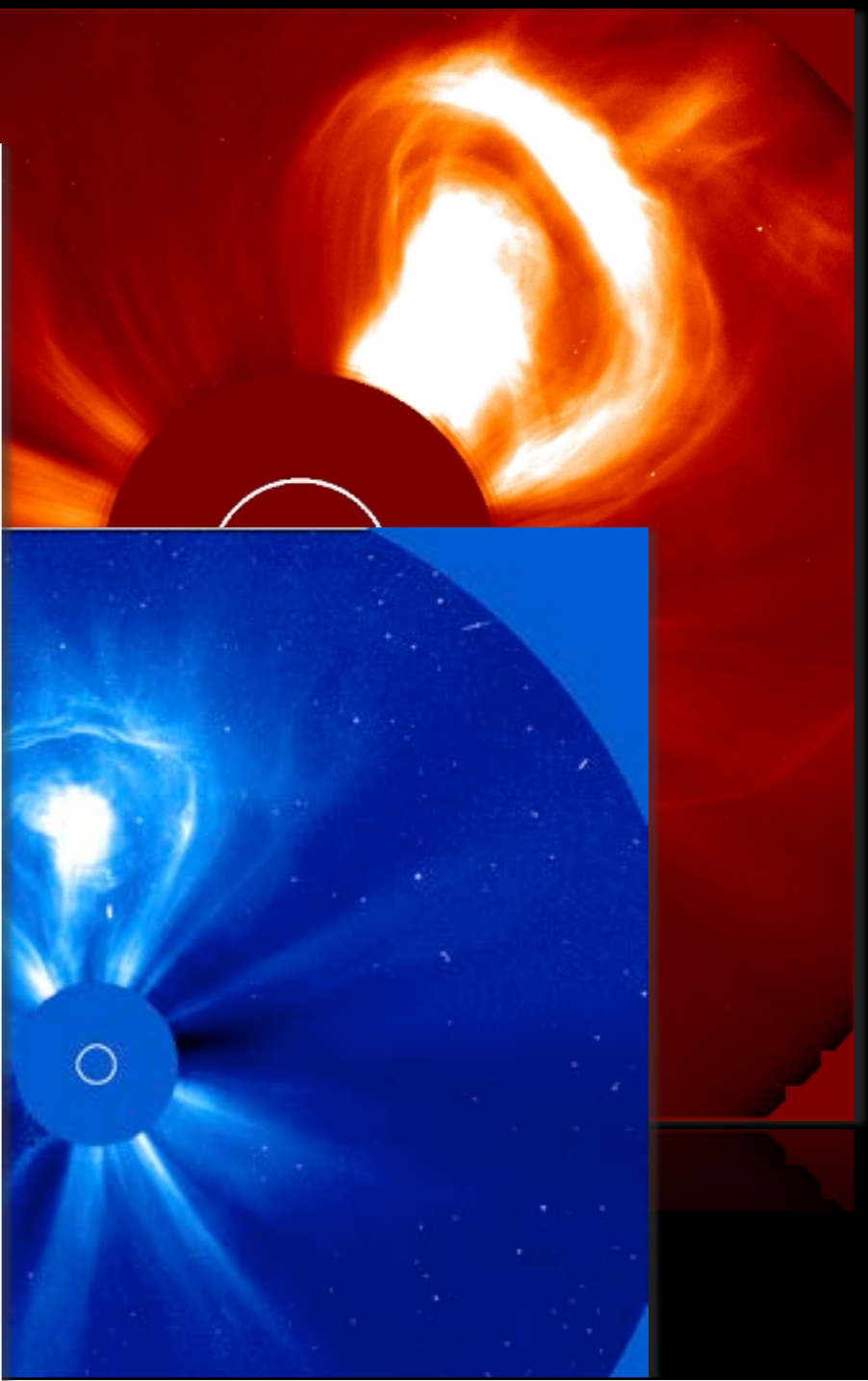
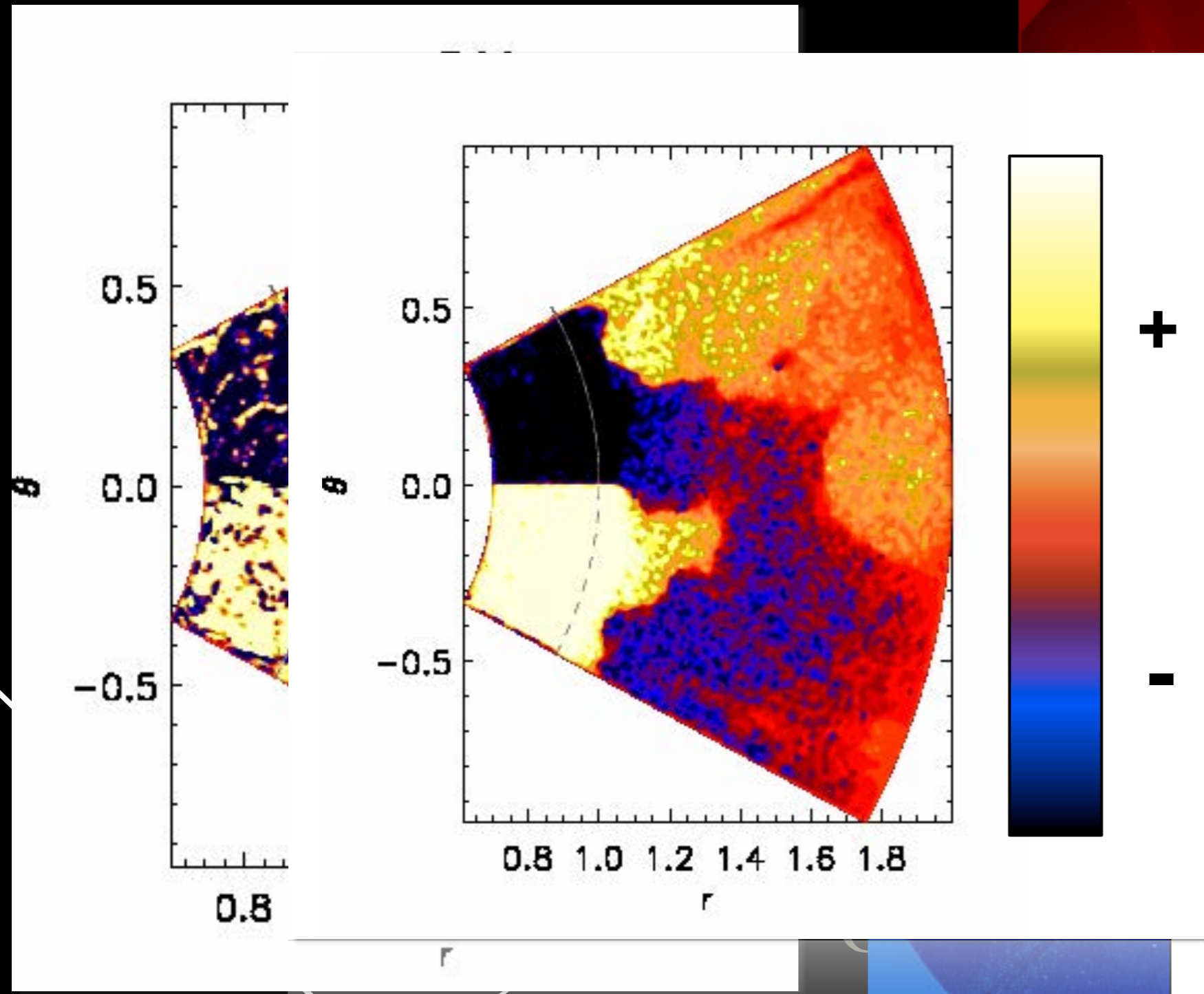
CME-like Ejections



Warnecke et al.
2011
(A&A 534, A11)

2000/02/27 07:42
54:70 15:50\0005

CME-like Ejections



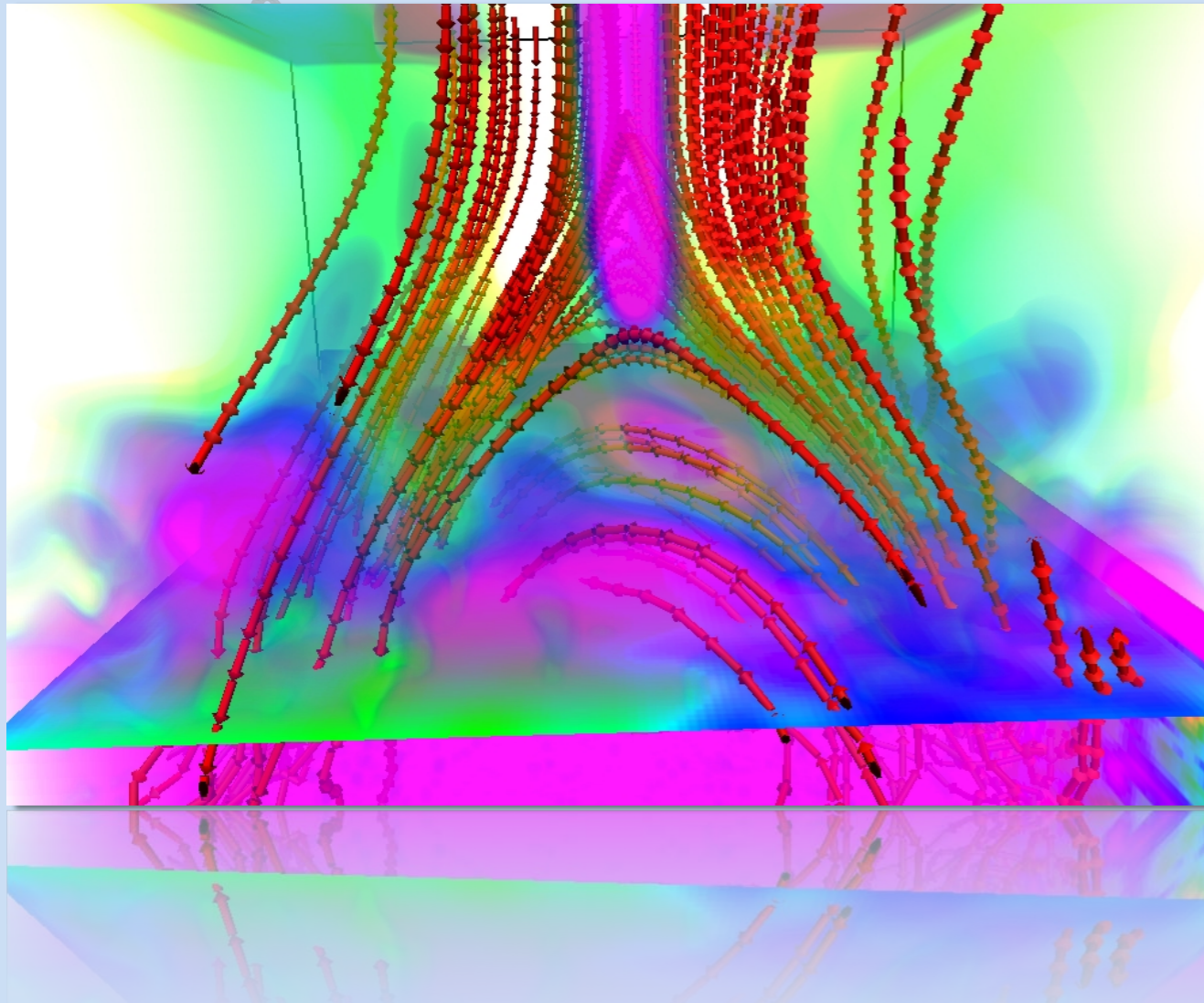
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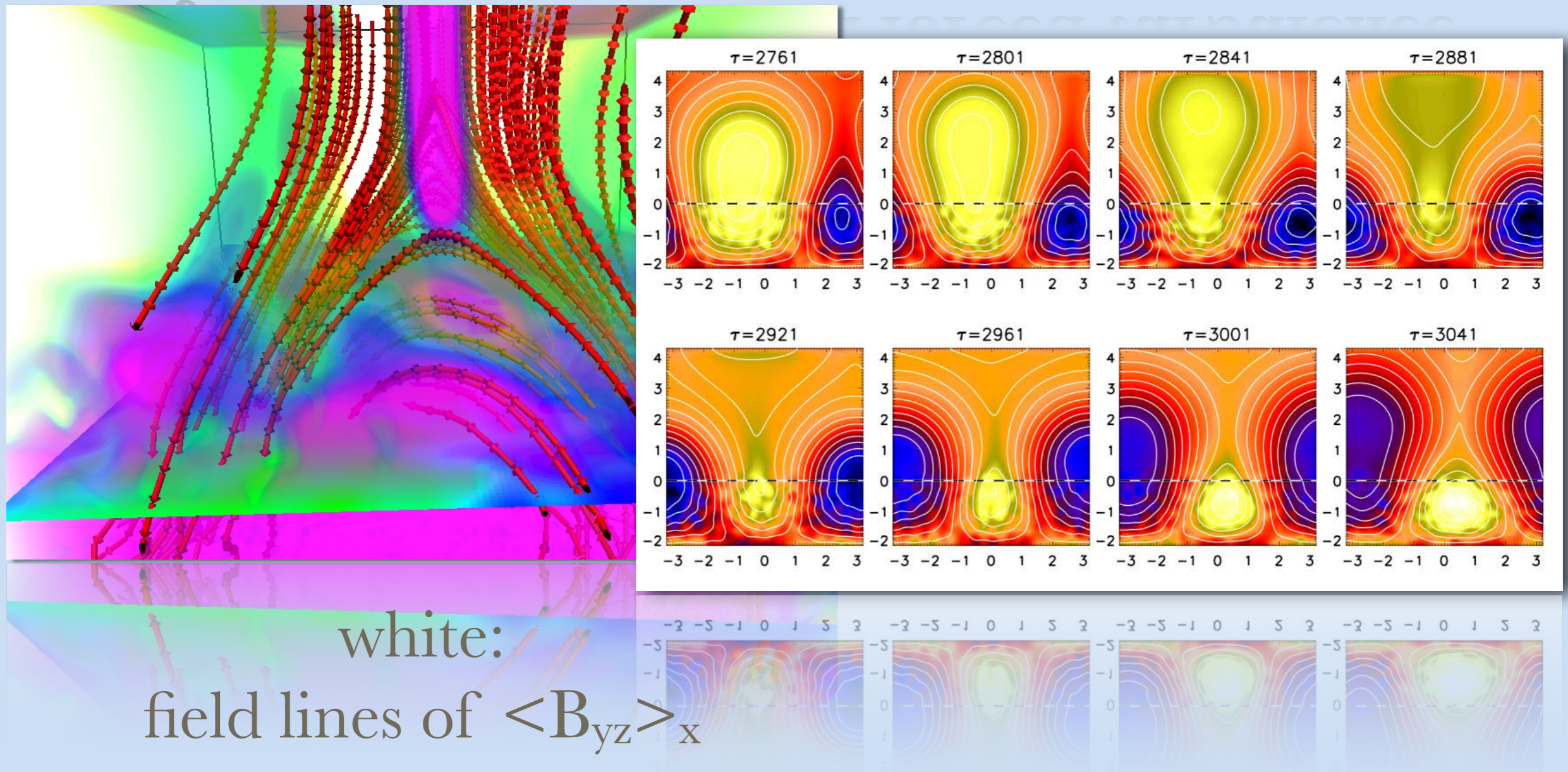
Ejections in Cartesian forced turbulence

filamentary structures in Cartesian forced turbulence

Ejections in Cartesian forced turbulence



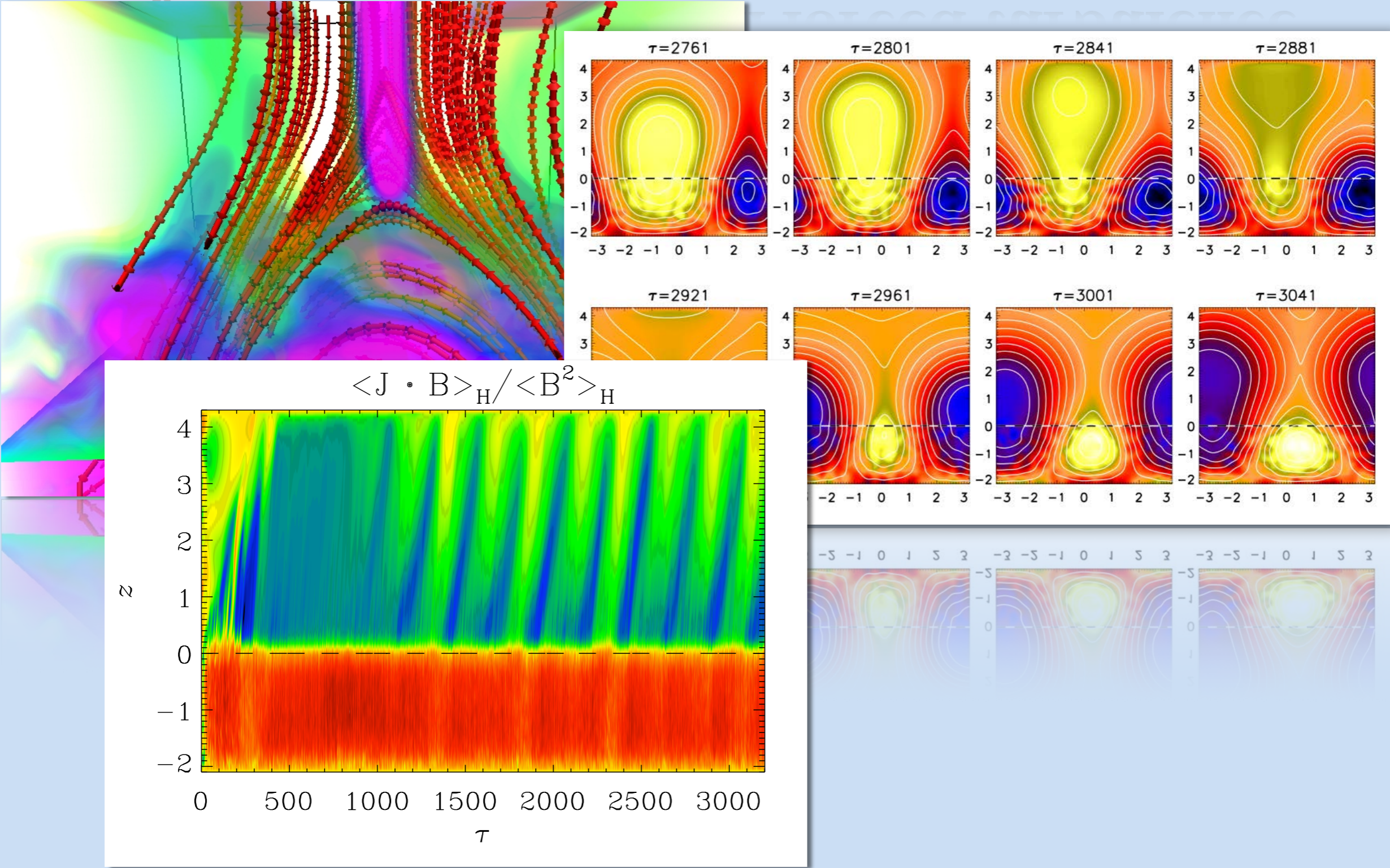
Ejections in Cartesian forced turbulence



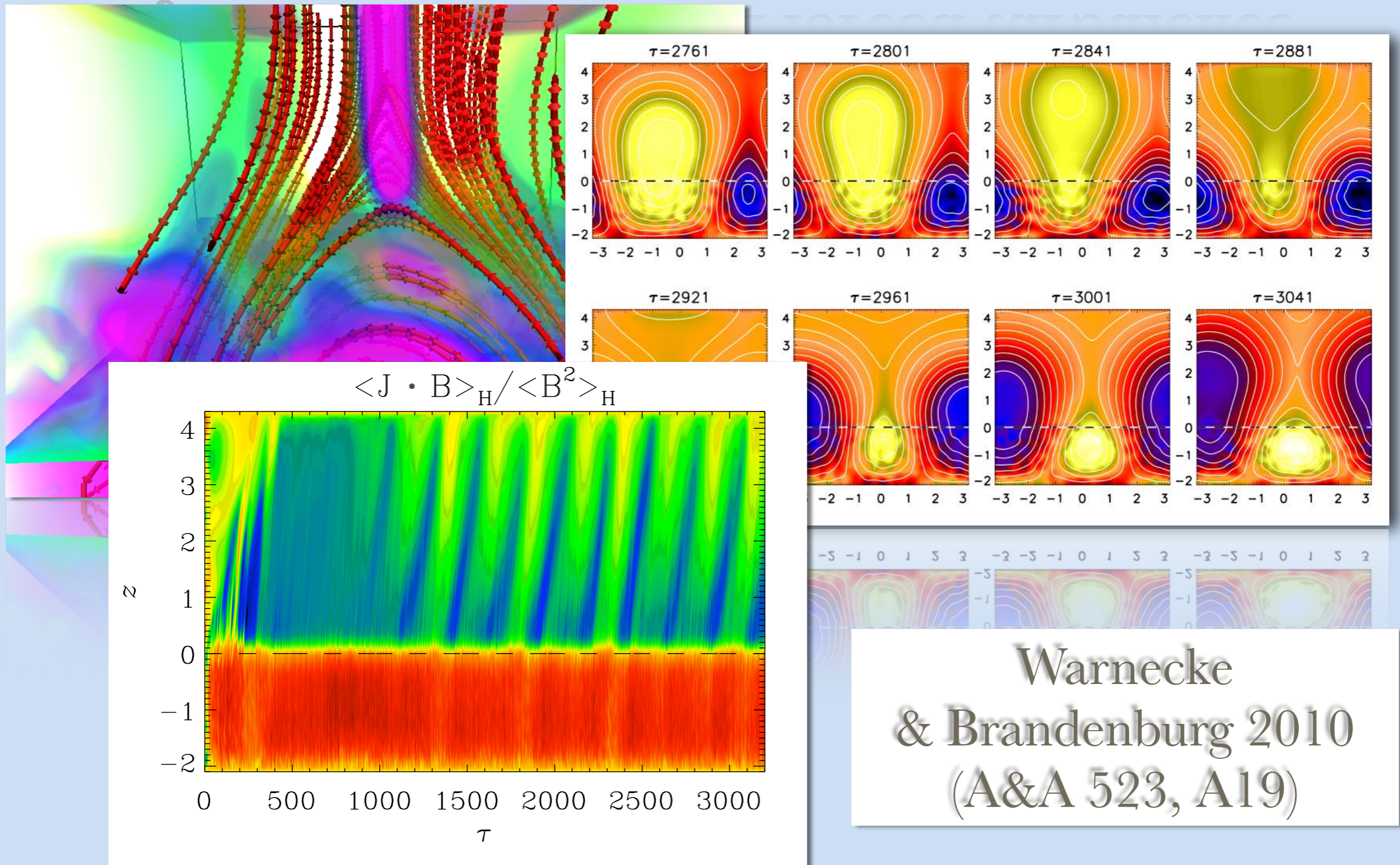
white:
field lines of $\langle B_{yz} \rangle_x$

color coded:
 $\langle B_x \rangle_x$

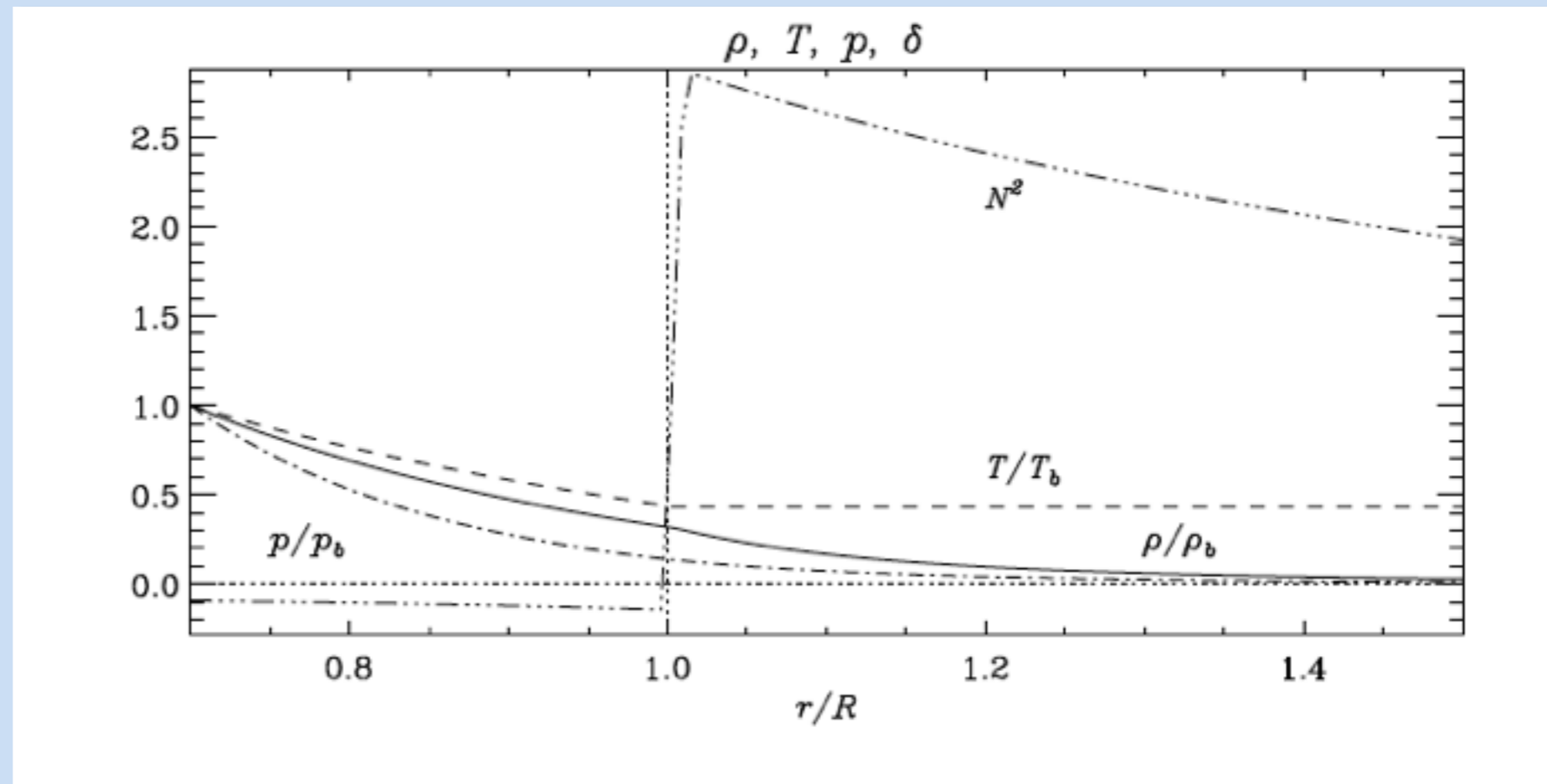
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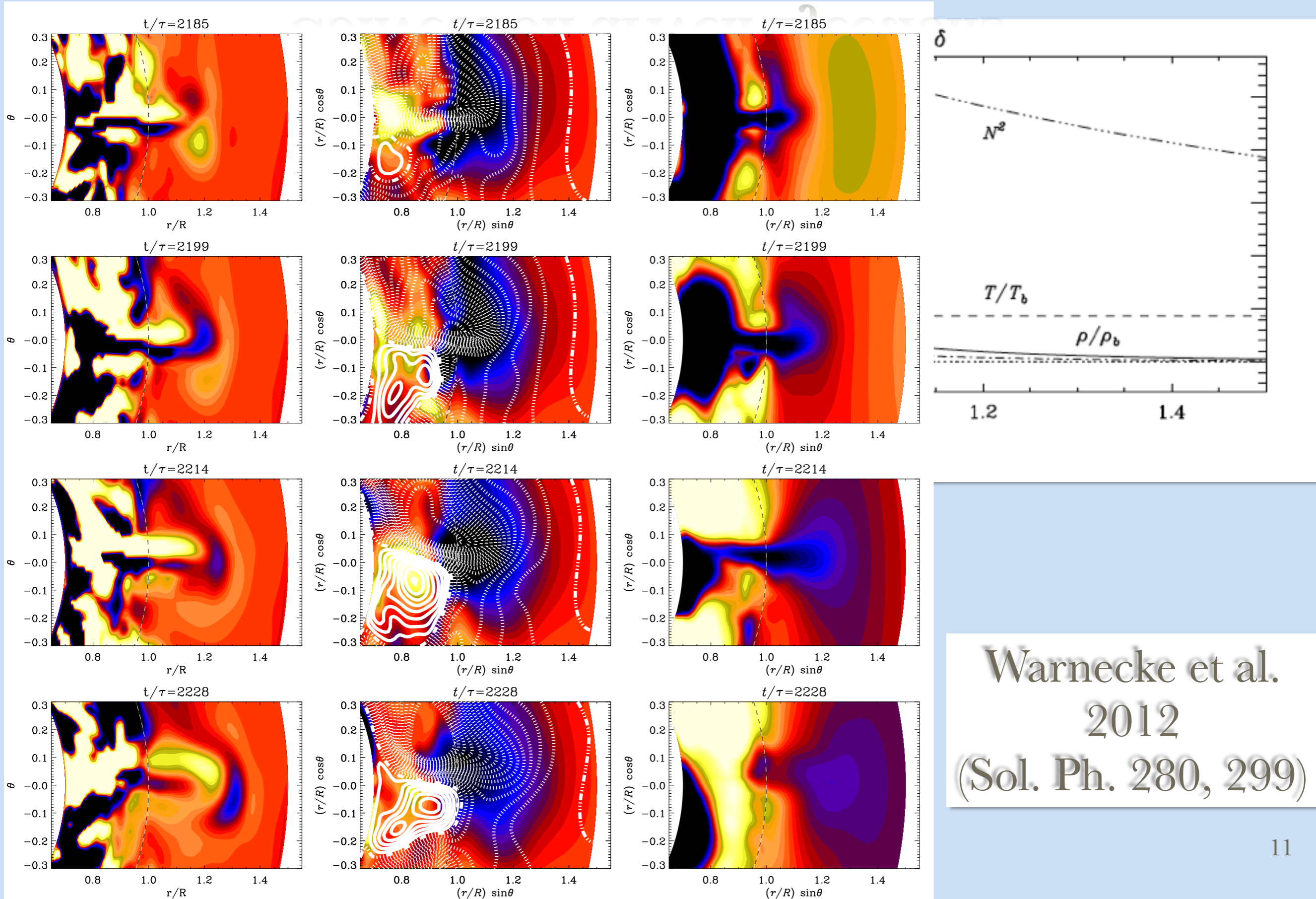


Warnecke
& Brandenburg 2010
(A&A 523, A19)



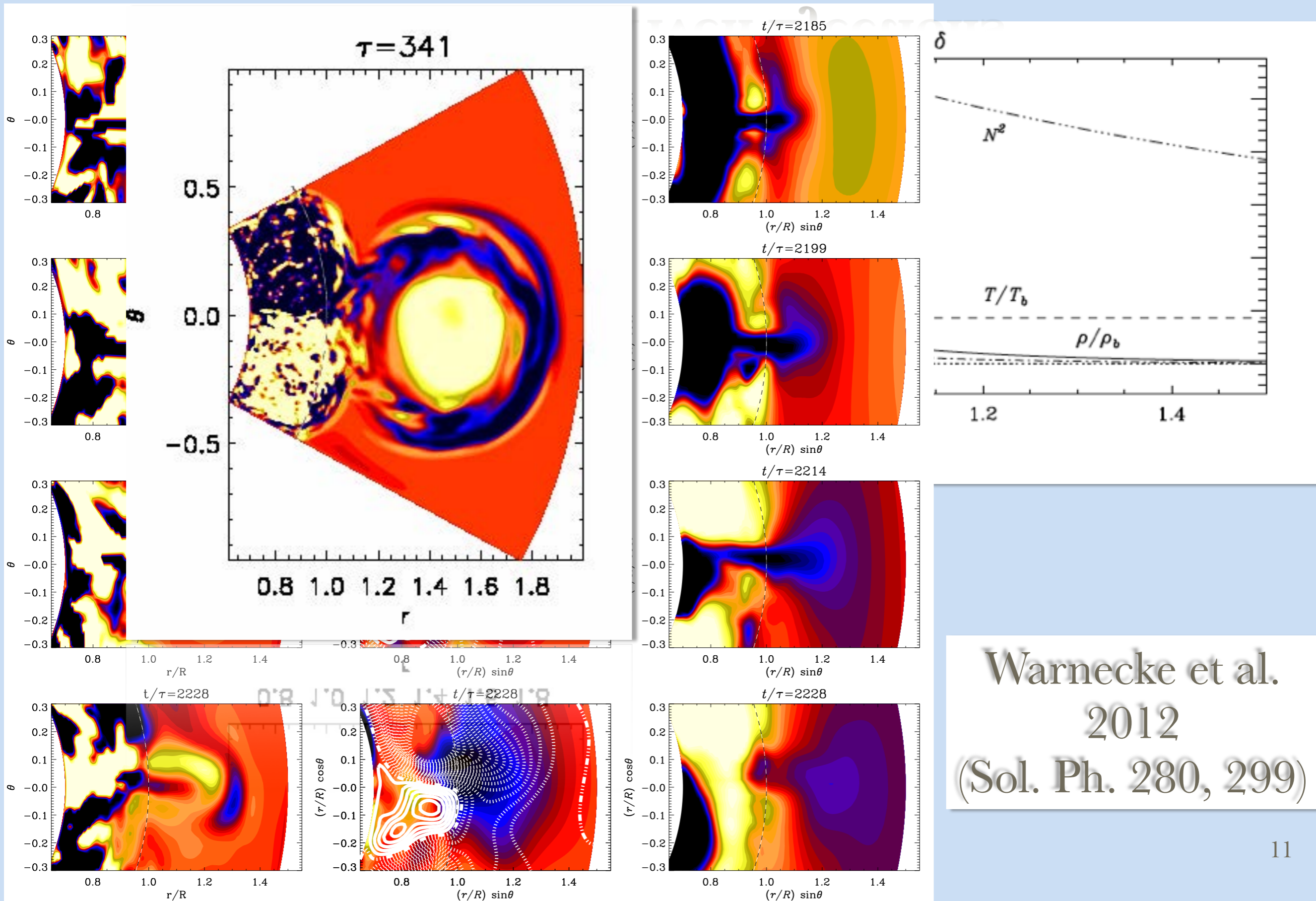
Self-consistent convection with a coronal layer

Convection driven Ejections



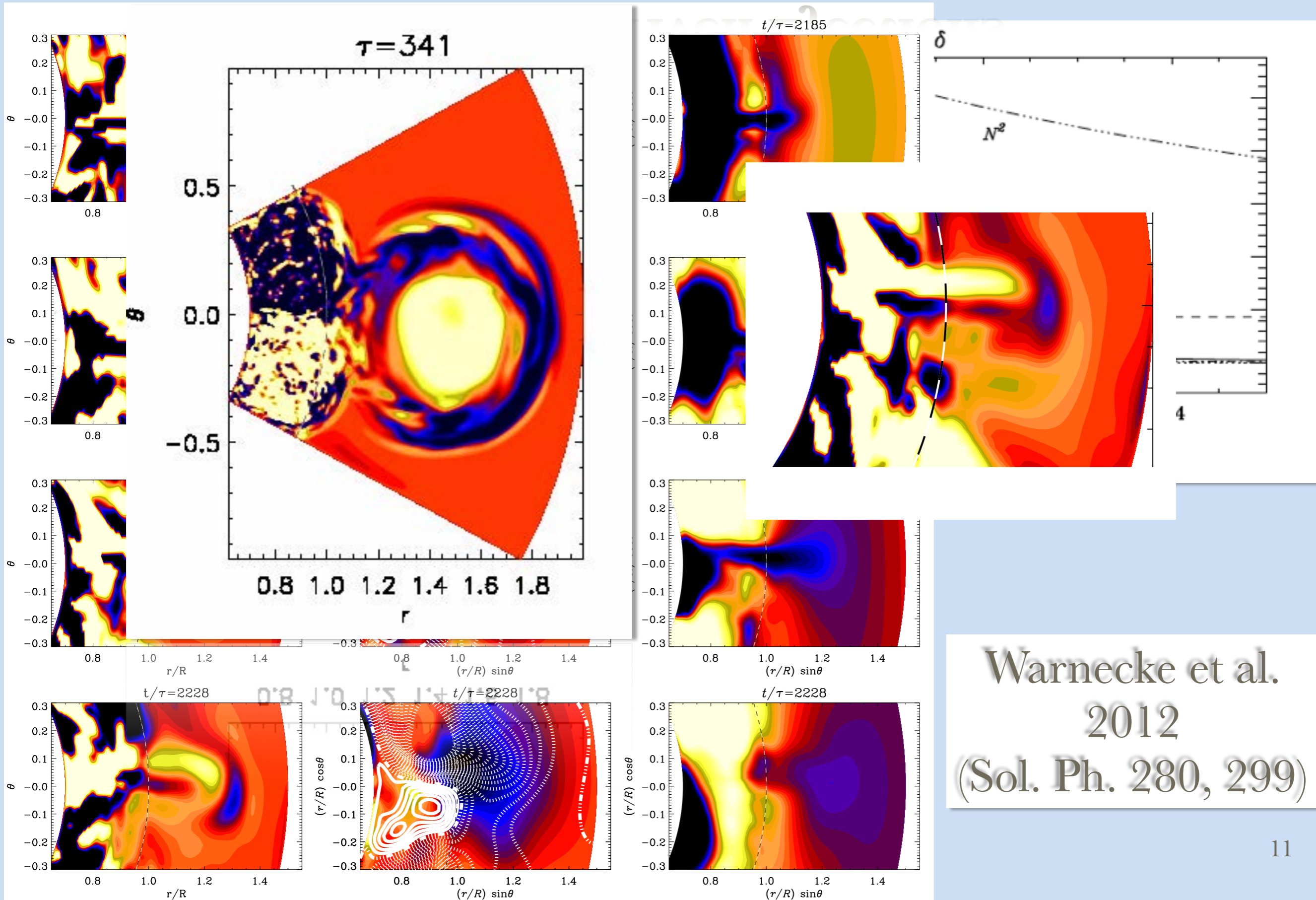
Warnecke et al.
2012
(Sol. Ph. 280, 299)

Convection driven Ejections



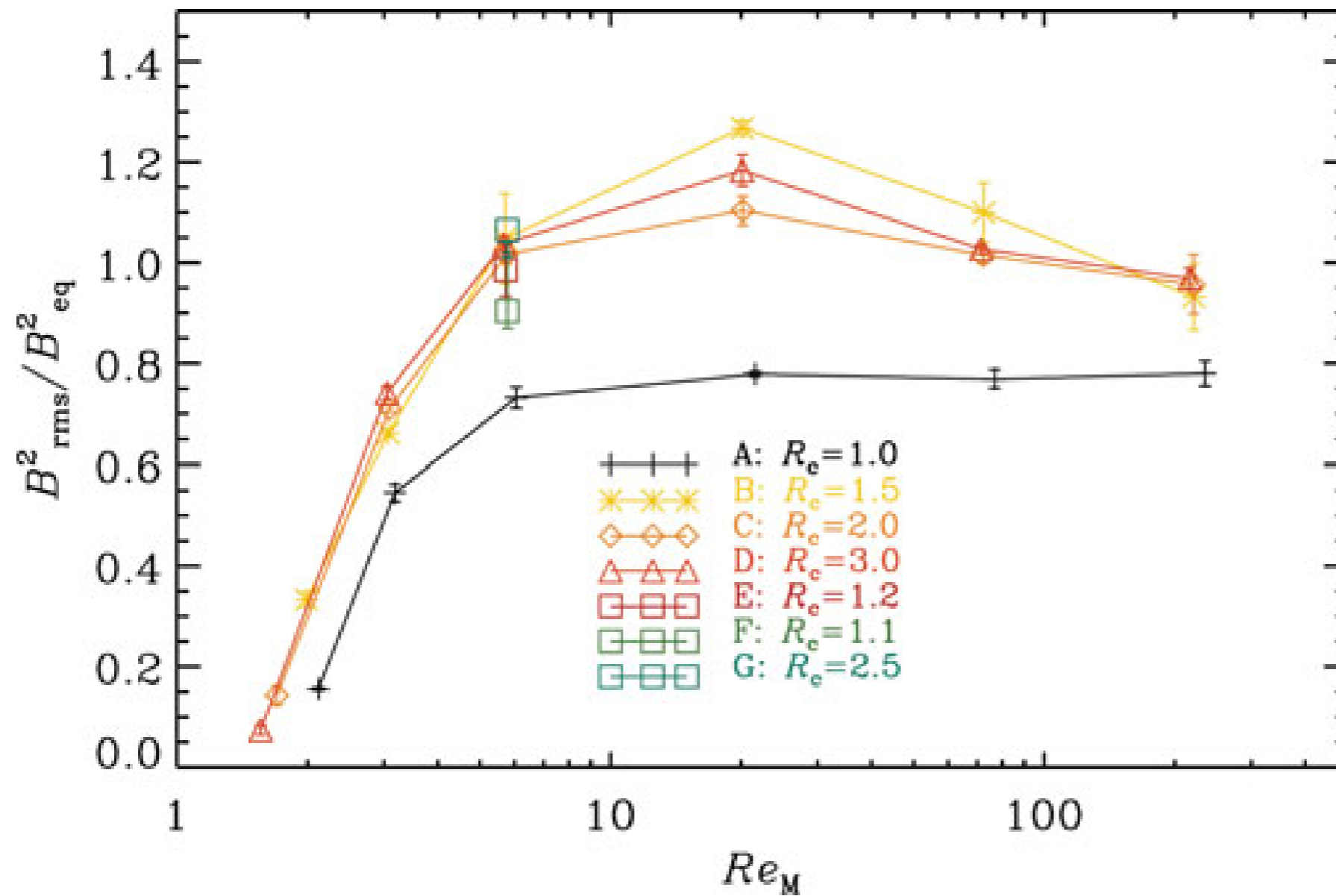
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Convection driven Ejections

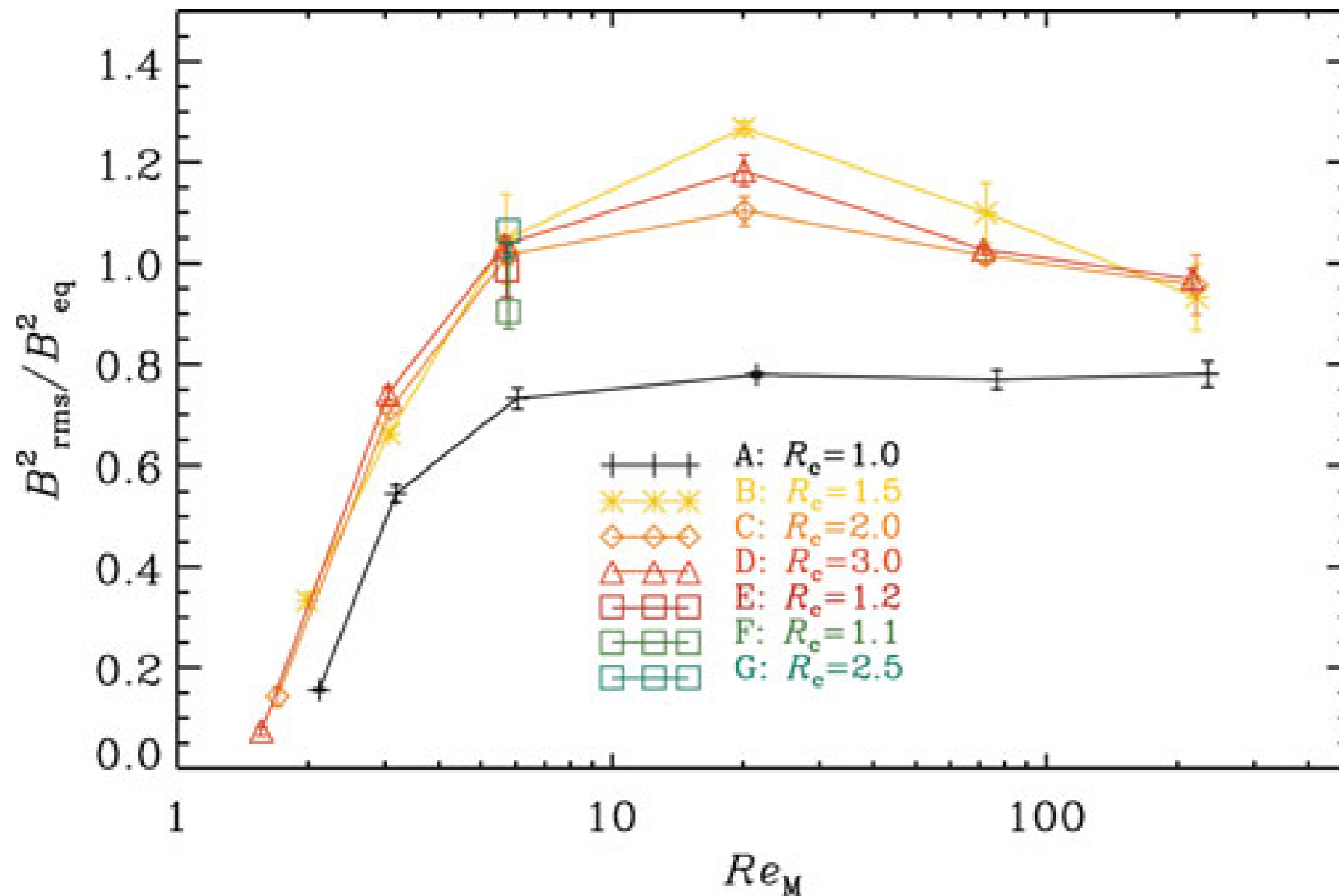


Warnecke et al.
2012
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Corona supports dynamo action



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Warnecke et al., 2014
IAU proceeding

Coronal model driven by emerging flux simulation

flux-emergence simulation

from / similar to Cheung et al (2010) ApJ 720, 233

– flux rope rises from bottom
and breaks through surface

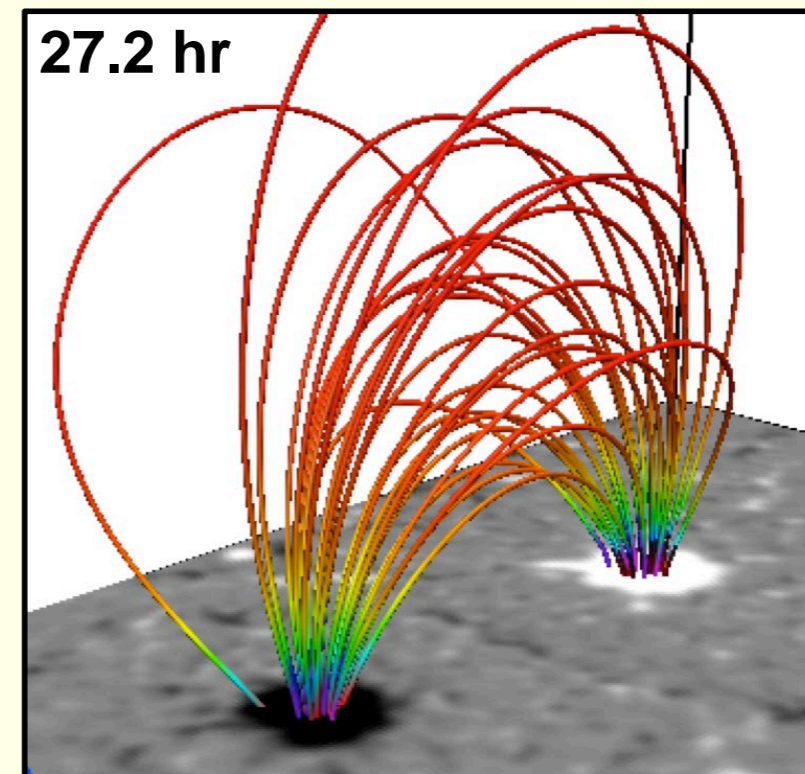
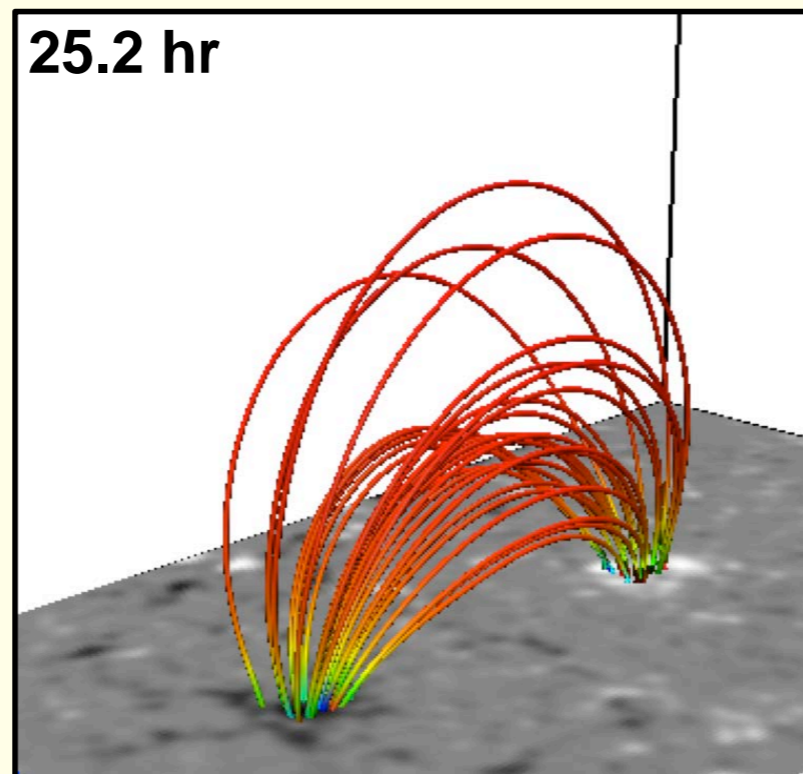
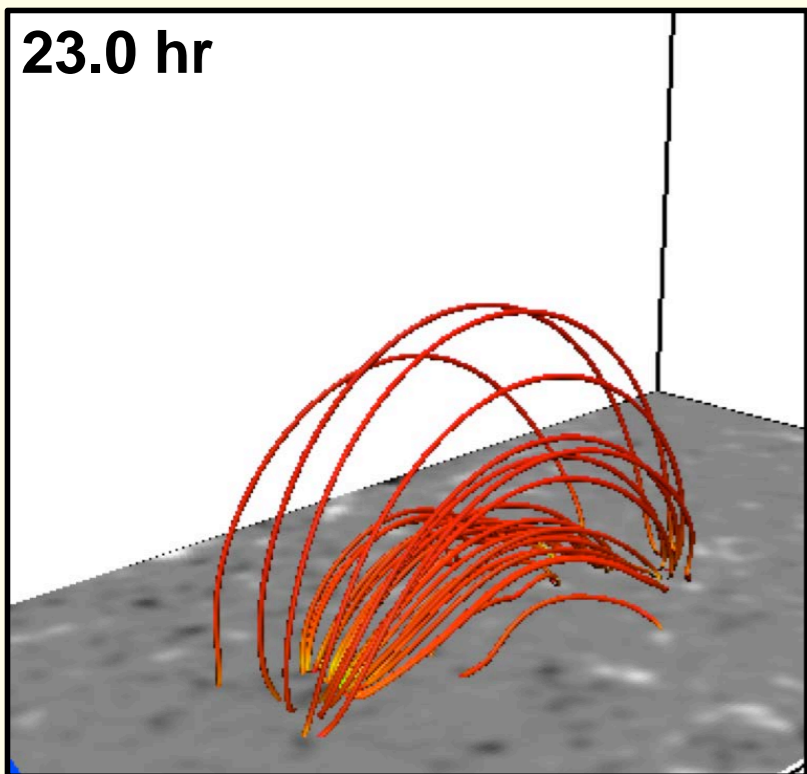
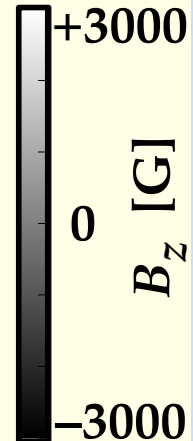
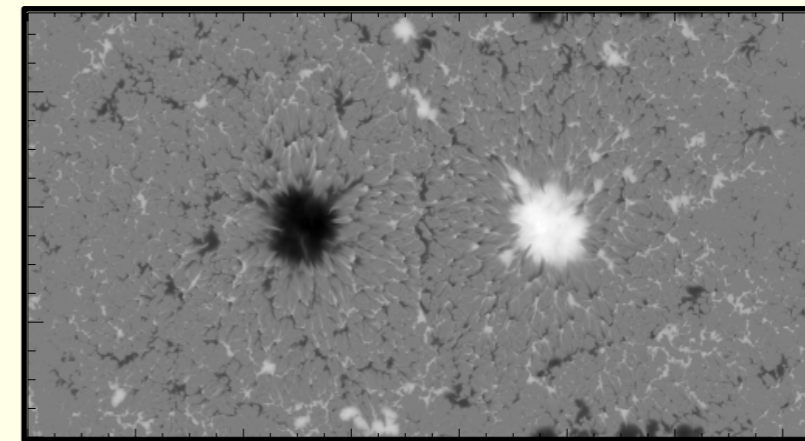
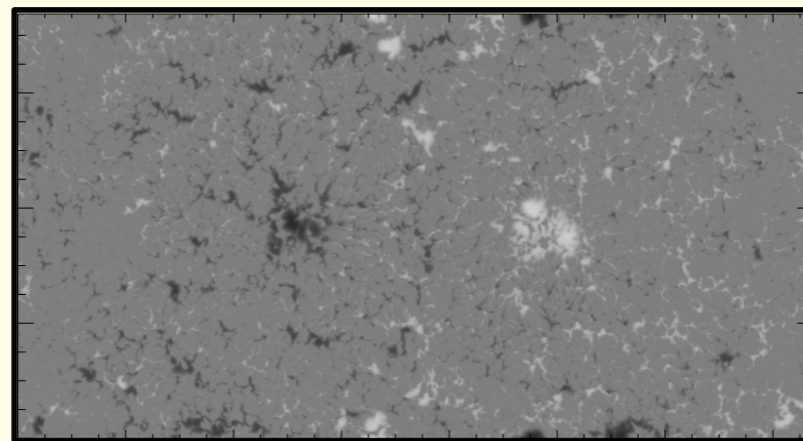
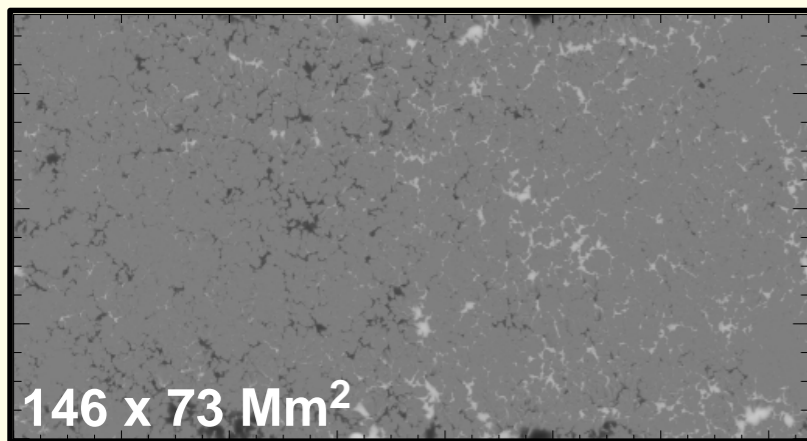
→ pair of sunspots

coronal simulation

– use photospheric layer (T, ρ, v, B)
as time-dependent lower boundary

→ magnetic field expands

→ coronal loops form

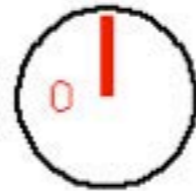
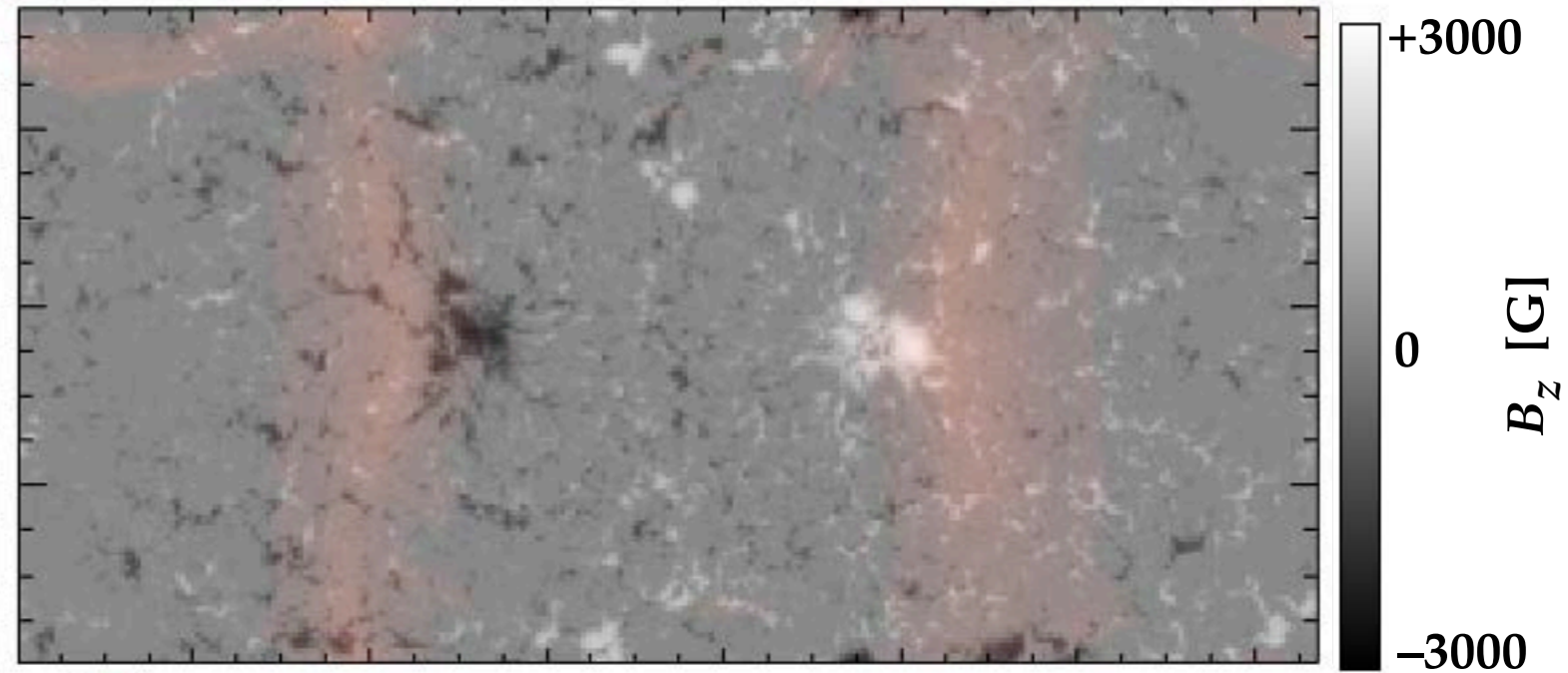


Chen, Bingert, Peter, Cheung (2013)

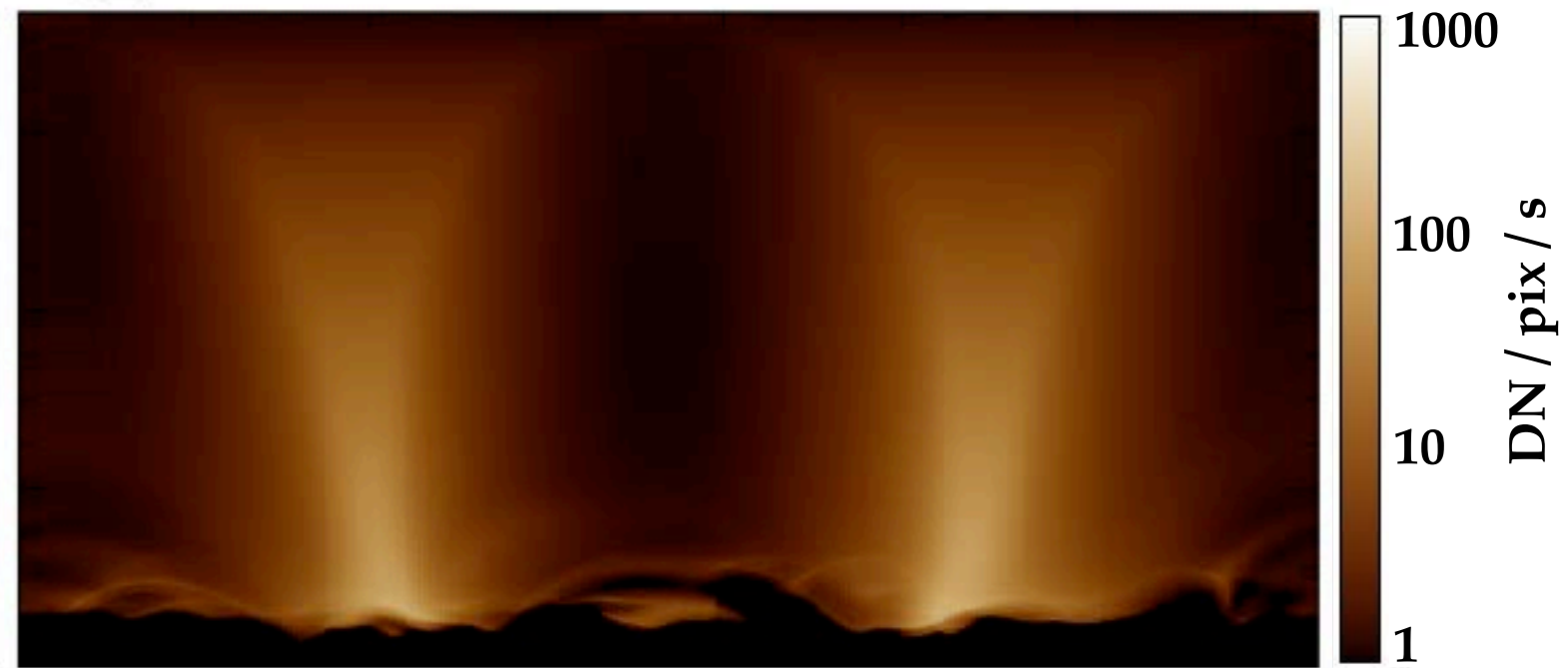
Coronal model driven by emerging flux simulation

synthesized coronal emission (1.5×10^6 K)

view from top: B_{vert} @ bottom + AIA 193 Å



view from side: AIA 193 Å



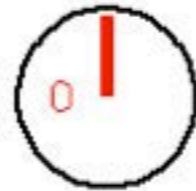
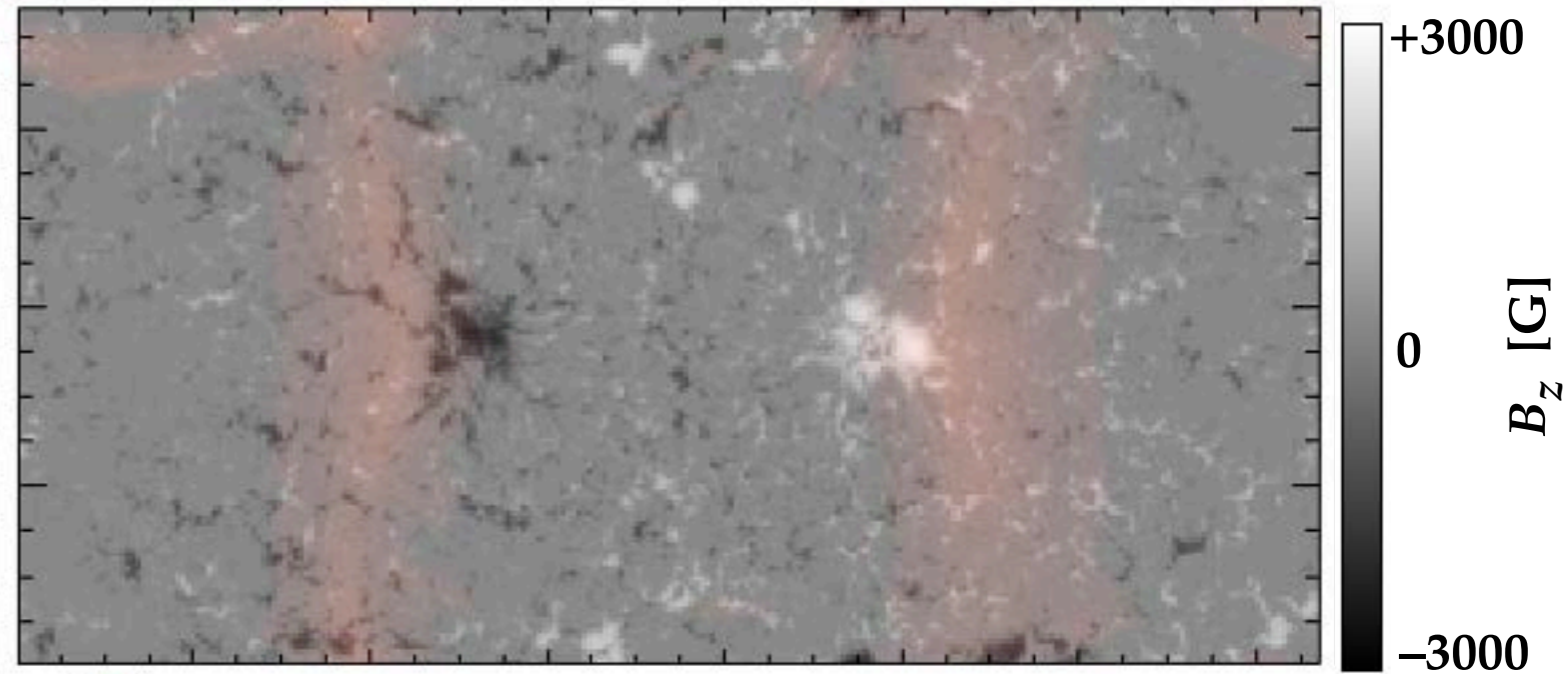
34 min
out of 10 hrs

- ▶ loops form at different places at different times
- ▶ loop footpoints are in sunspot periphery (penumbra)

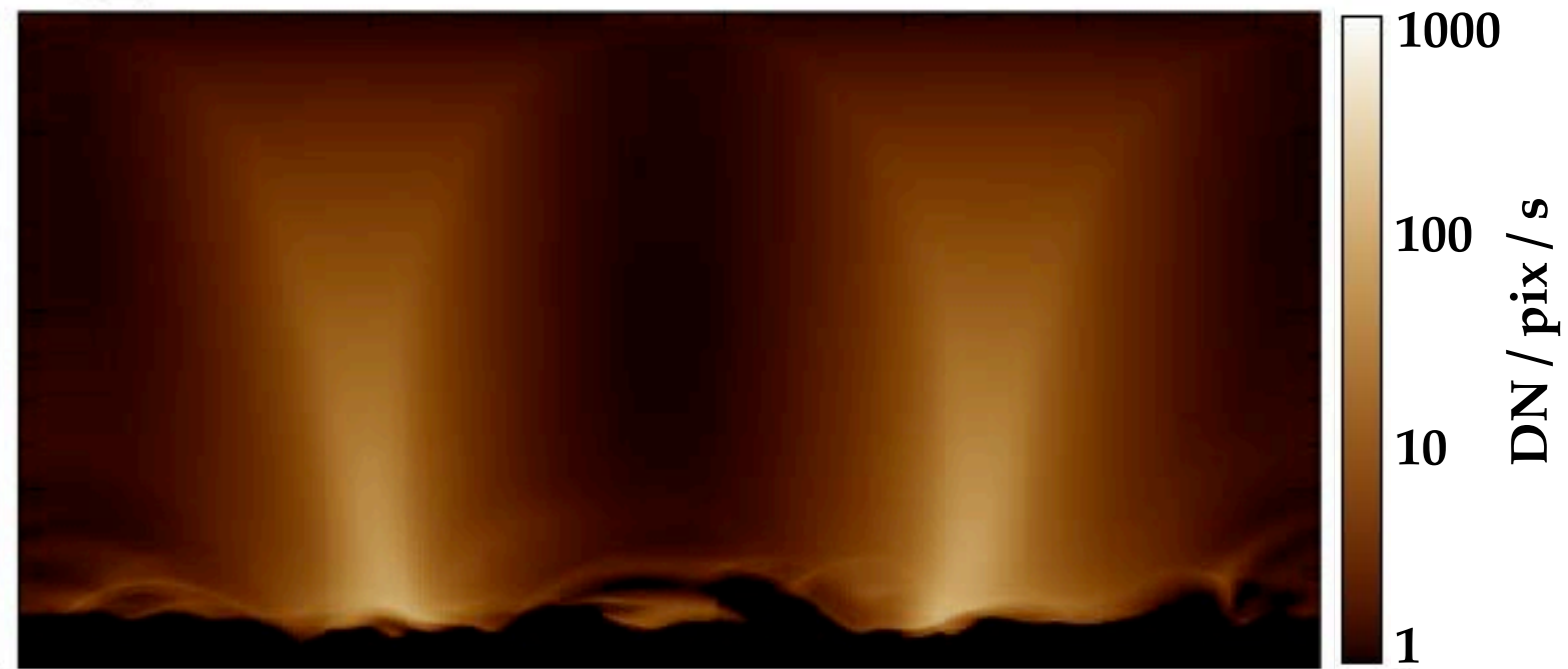
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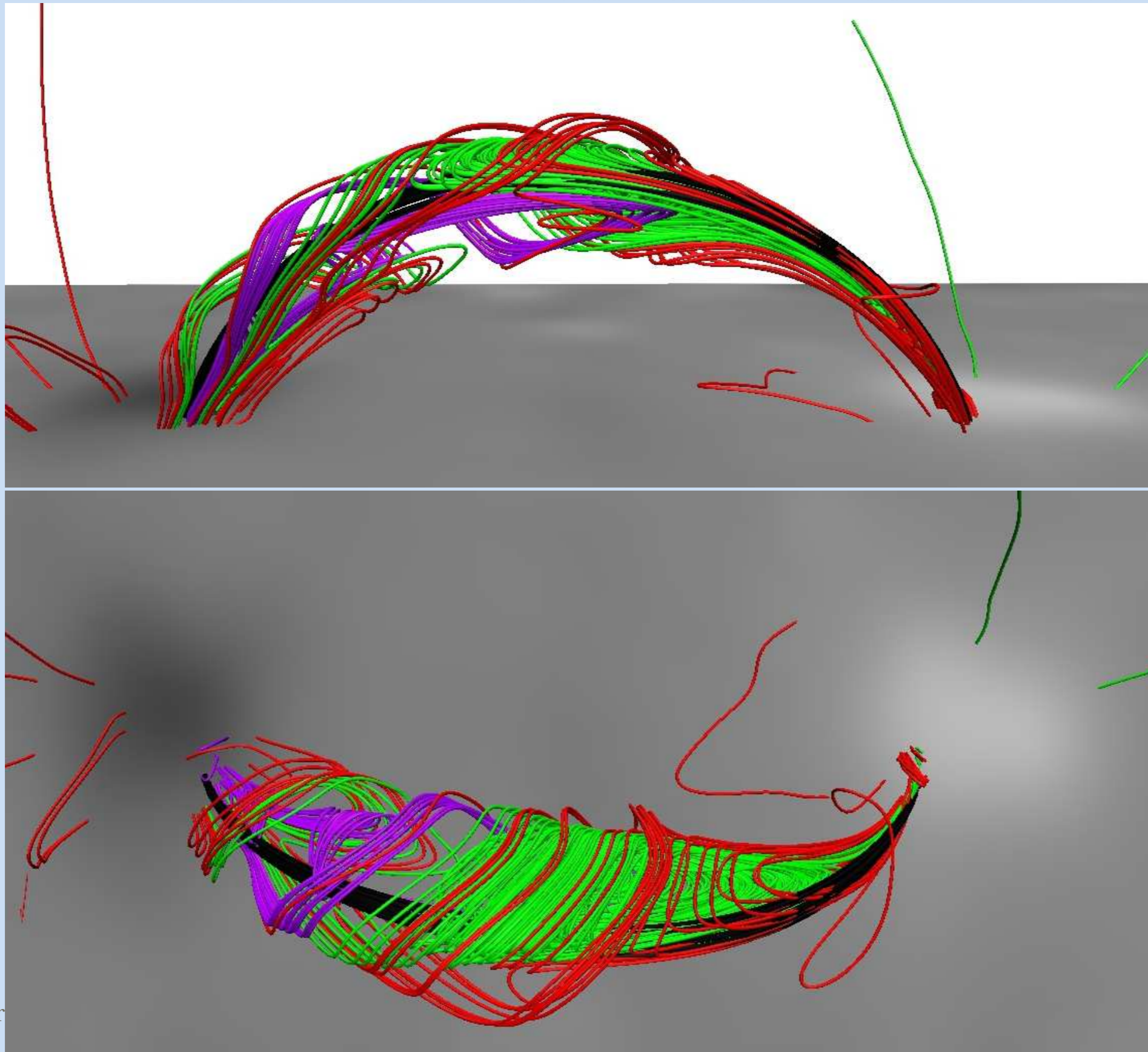
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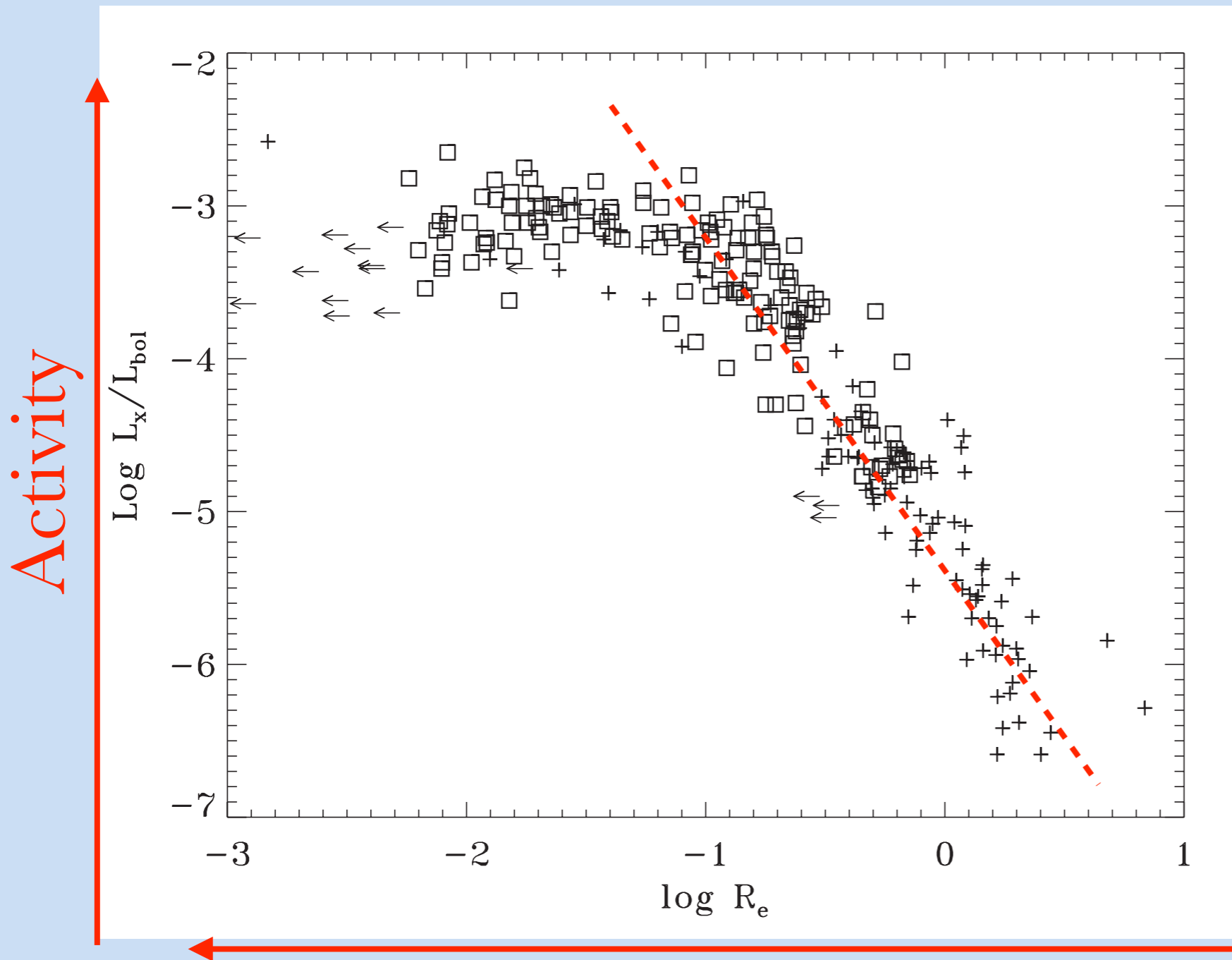
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Helical currents in coronal loops



Rotation Activity Relation

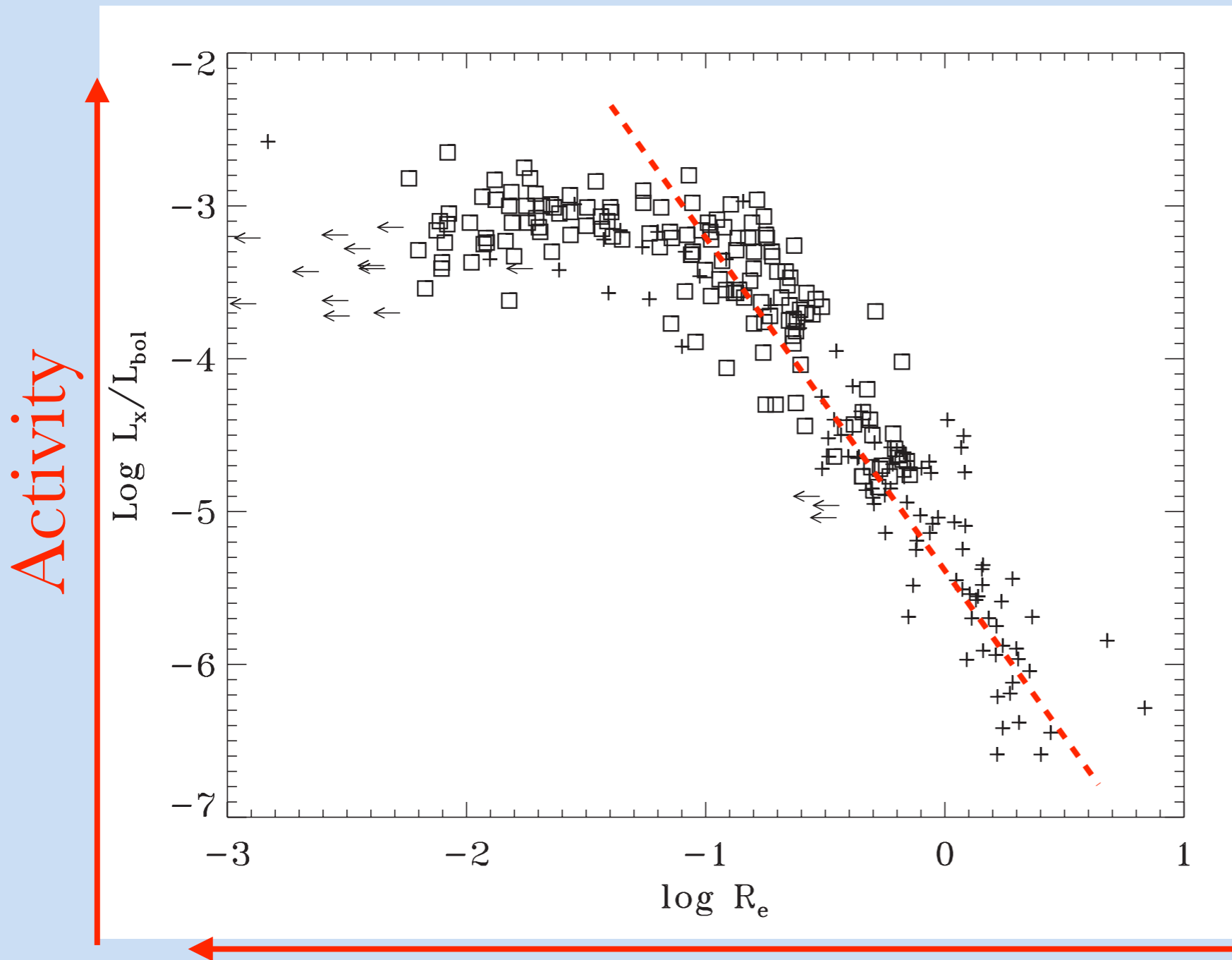


Pizzolato et al. 2003

Rotation

Rotation Activity Relation

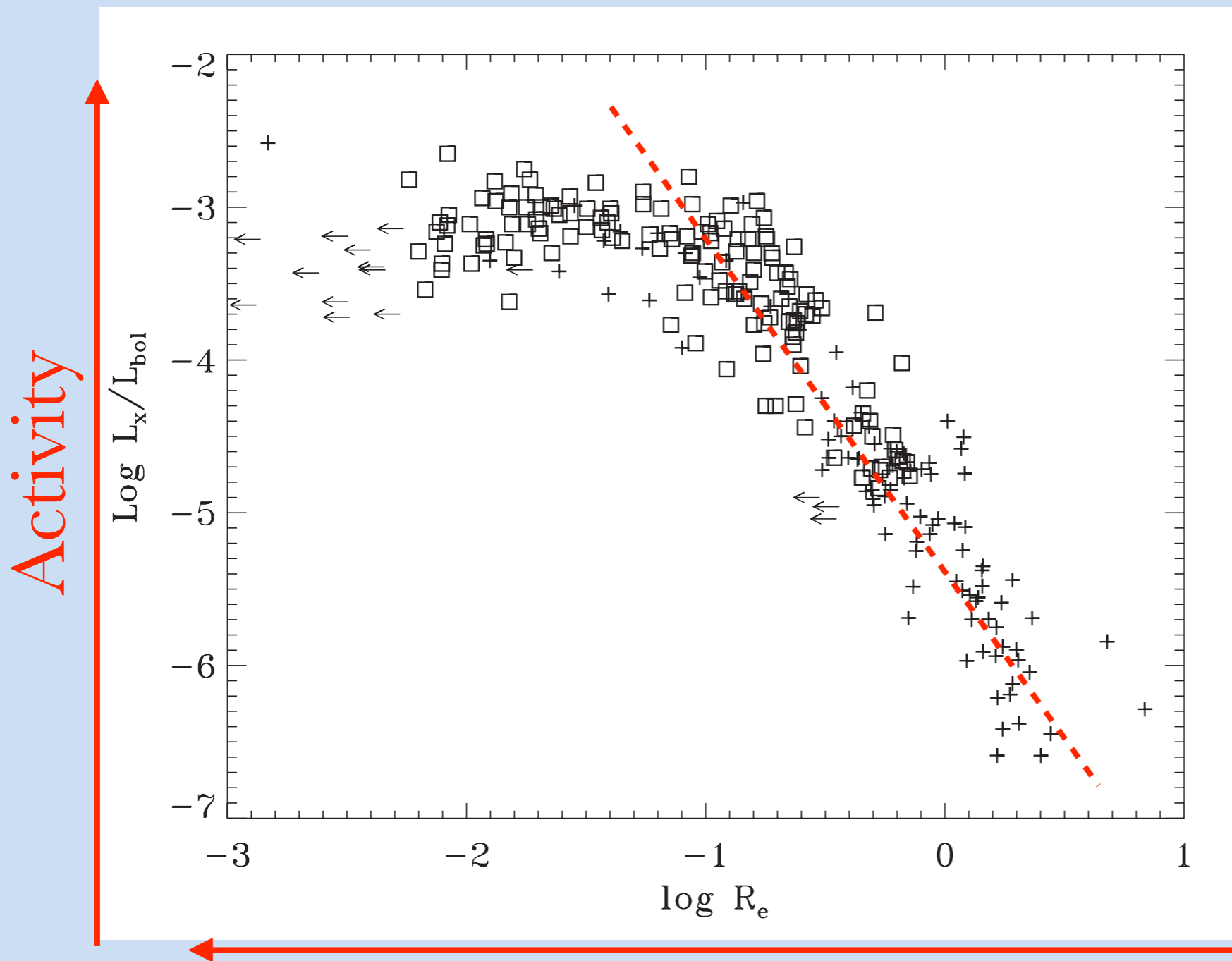
$$\nabla\Omega = \text{const}$$



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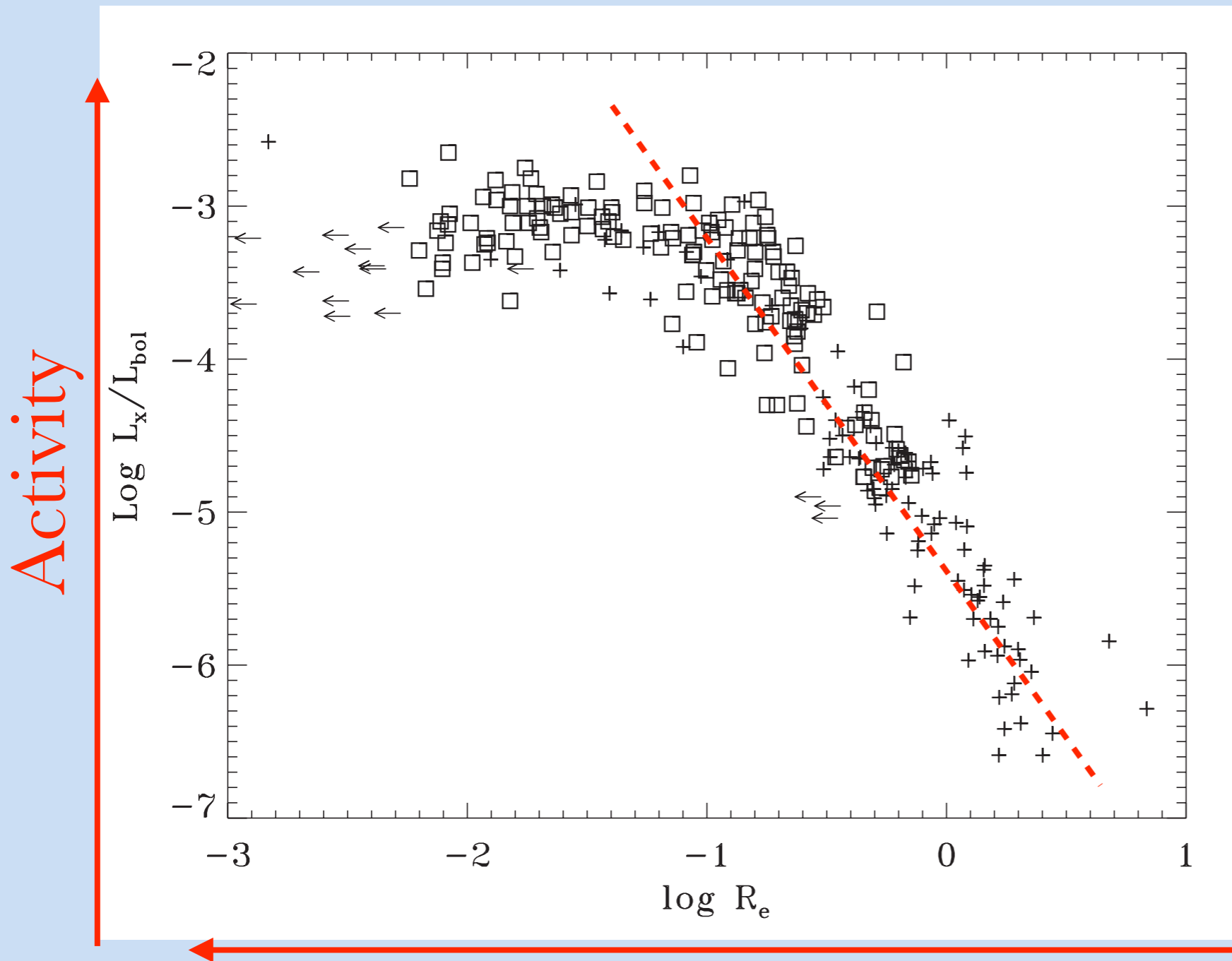
Rotation

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Pouquet et al. 1976

Rotation Activity Relation



Pizzolato et al. 2003

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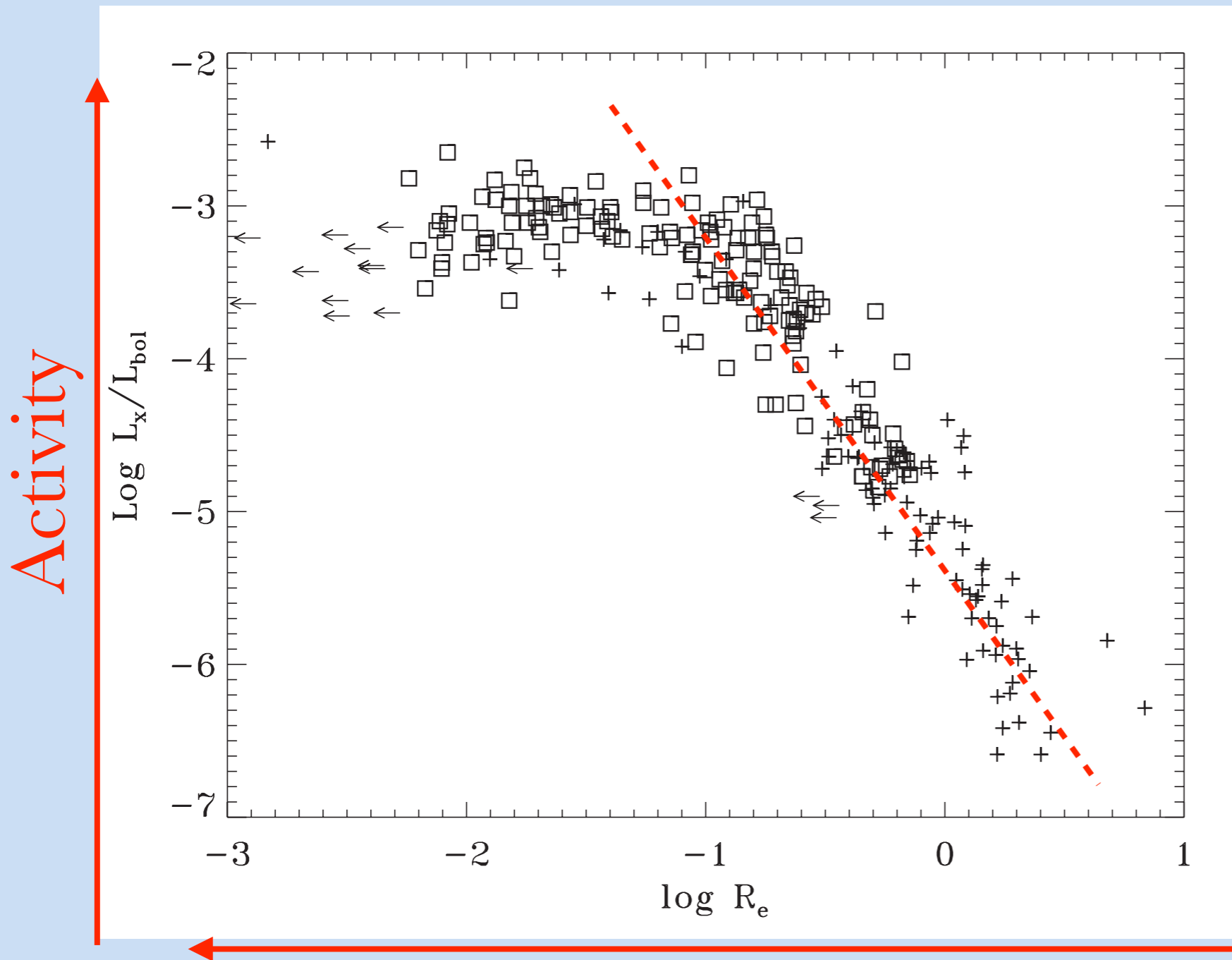
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Pouquet et al. 1976

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Rotation Activity Relation



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Pouquet et al. 1976

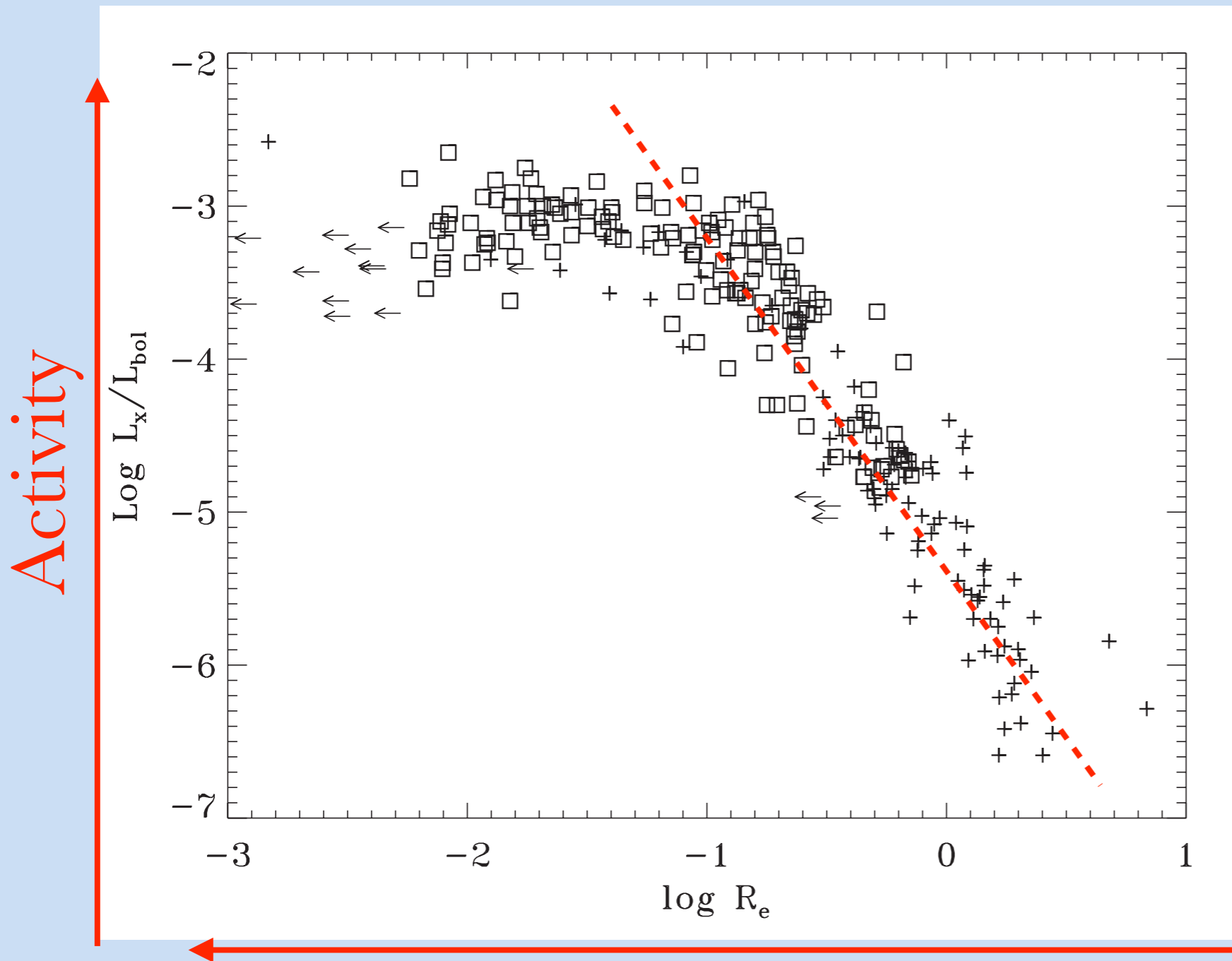
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Pizzolato et al. 2003

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Pouquet et al. 1976

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$$\text{Act.} \approx j \cdot b \approx \omega \cdot u$$

Conclusions

CONCLUSIONS

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- Helical dynamos can produce ejection of current helicity self-consistently.
- Simplified corona and ejections supports dynamo action.
- Helicity important for coronal heating.
- Rotation might lead to enhanced activity due to helicity.